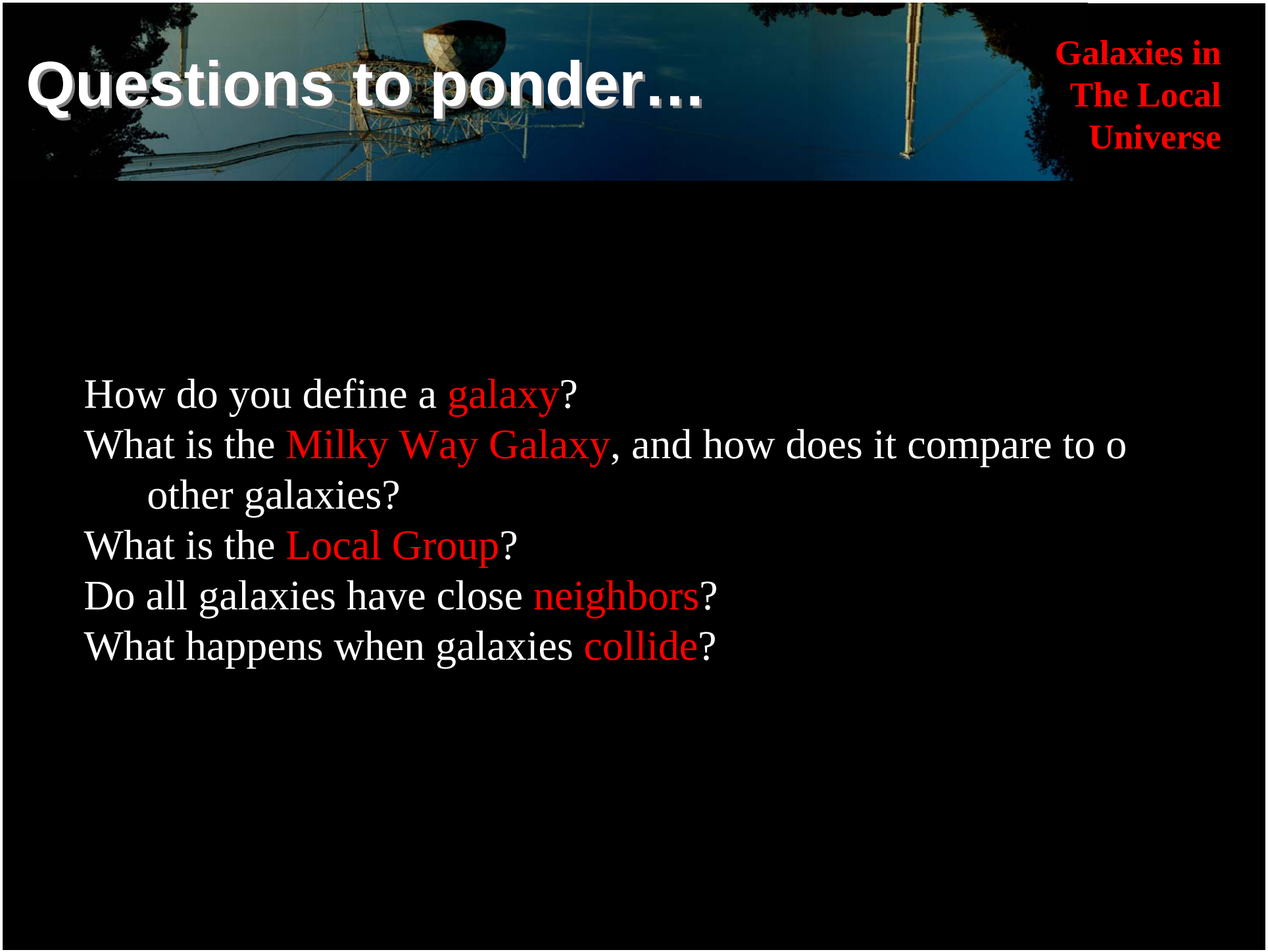
Galaxies in The Local Universe

Cornell University



Questions to ponder…

How do you define a galaxy?
What is the Milky Way Galaxy, and how does it compare to o
other galaxies?
What is the Local Group?
Do all galaxies have close neighbors?
What happens when galaxies collide?

A simple calculation: Redshift

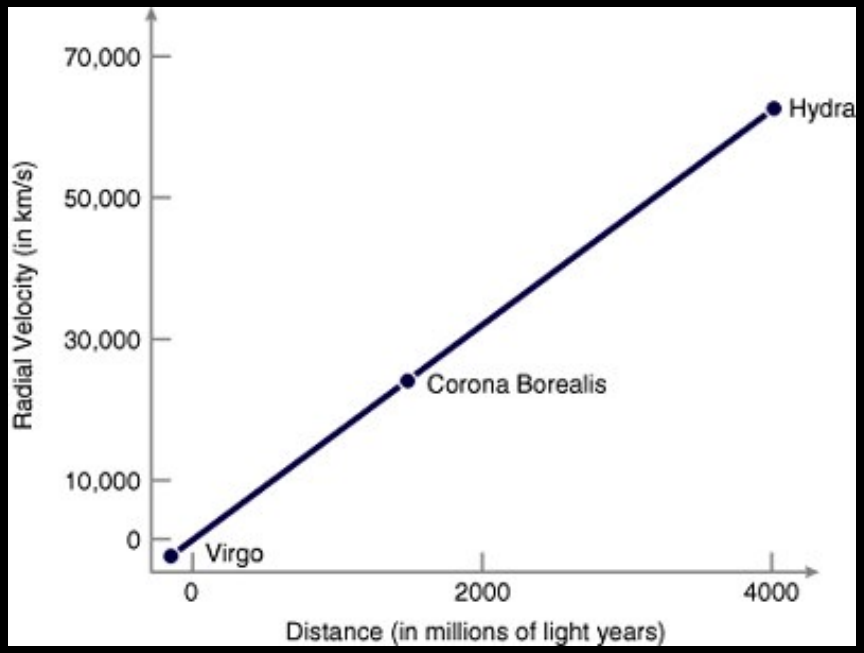
$$z = \frac{\lambda_{obs} - \lambda_0}{\lambda_0} = \frac{f_0 - f_{obs}}{f_{obs}}$$

- Measure the shift in a spectral line – f_0 is the rest frequency (λ_0 the rest wavelength)
- Extragalactic objects often identified by their *cz* measurement.
- ALFALFA will cover *cz* = –2000 to 17000 km/s (out to 250 Mpc)

Expansion of the Universe

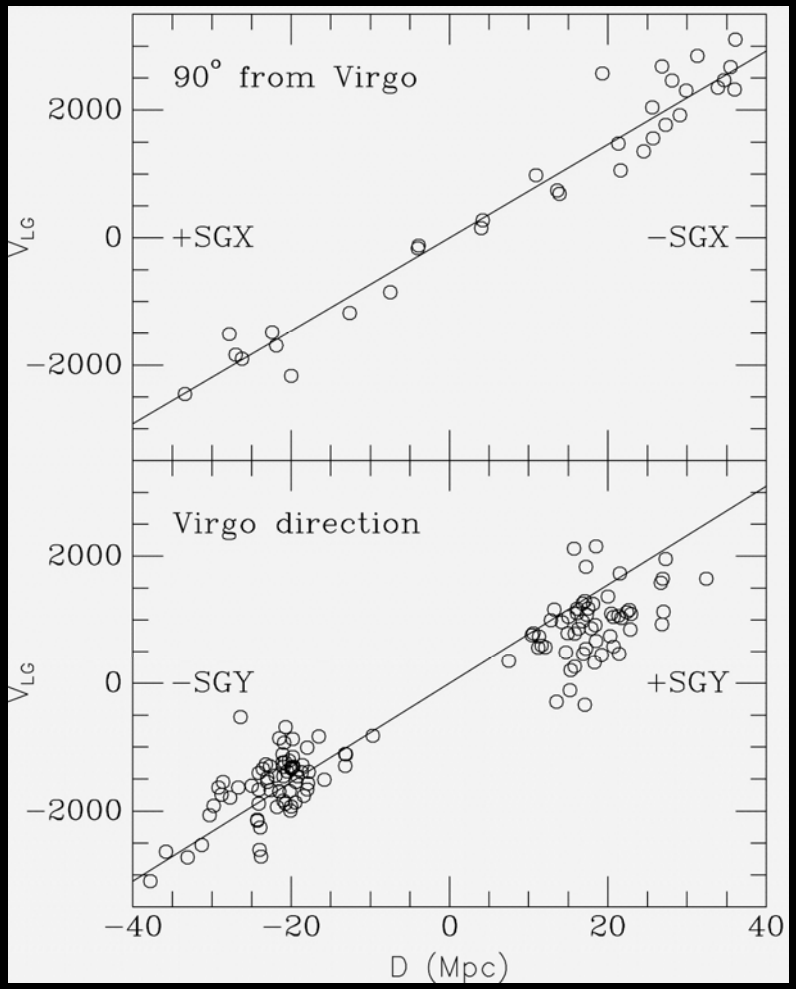
- Edwin Hubble showed the Universe was expanding!
- However, there are other factors to take into account in the local Universe – peculiar velocities! Deviations can be quite large depending on the galaxy, and whether it is part of a group or a field galaxy.

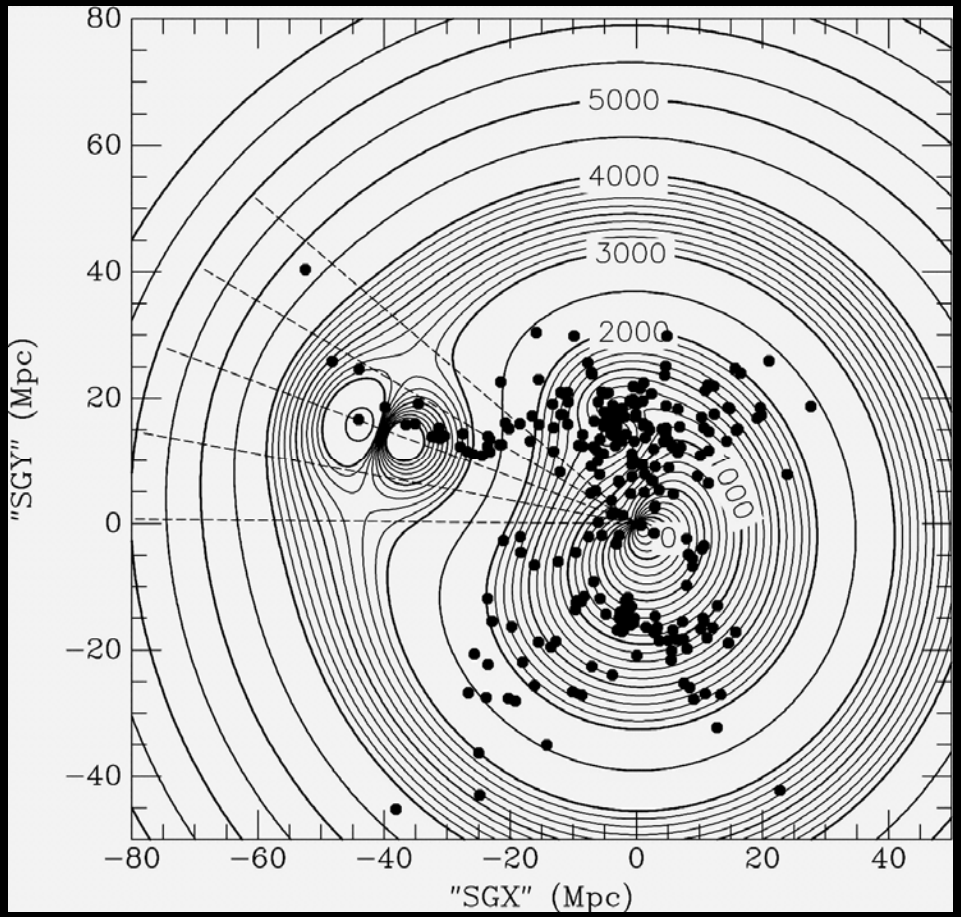
$$cz = H_0 d$$



Distances to nearby galaxies

Galaxies in The Local Universe



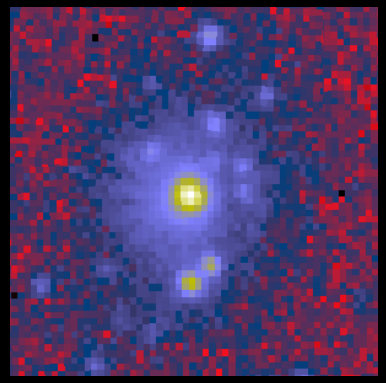
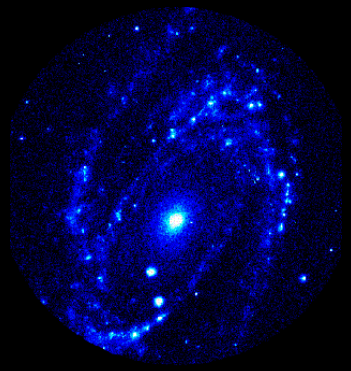


Tonry, et al. 2000


Galaxy Morphology

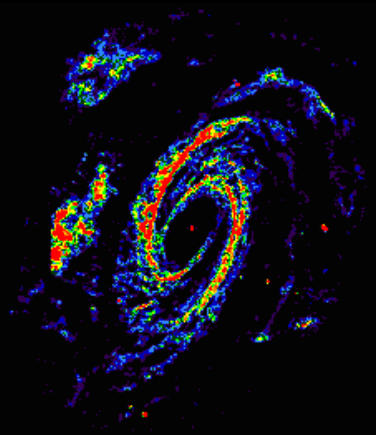
Galaxies across the spectrum

What do galaxies *look* like?

M81



cz = -34 km/s

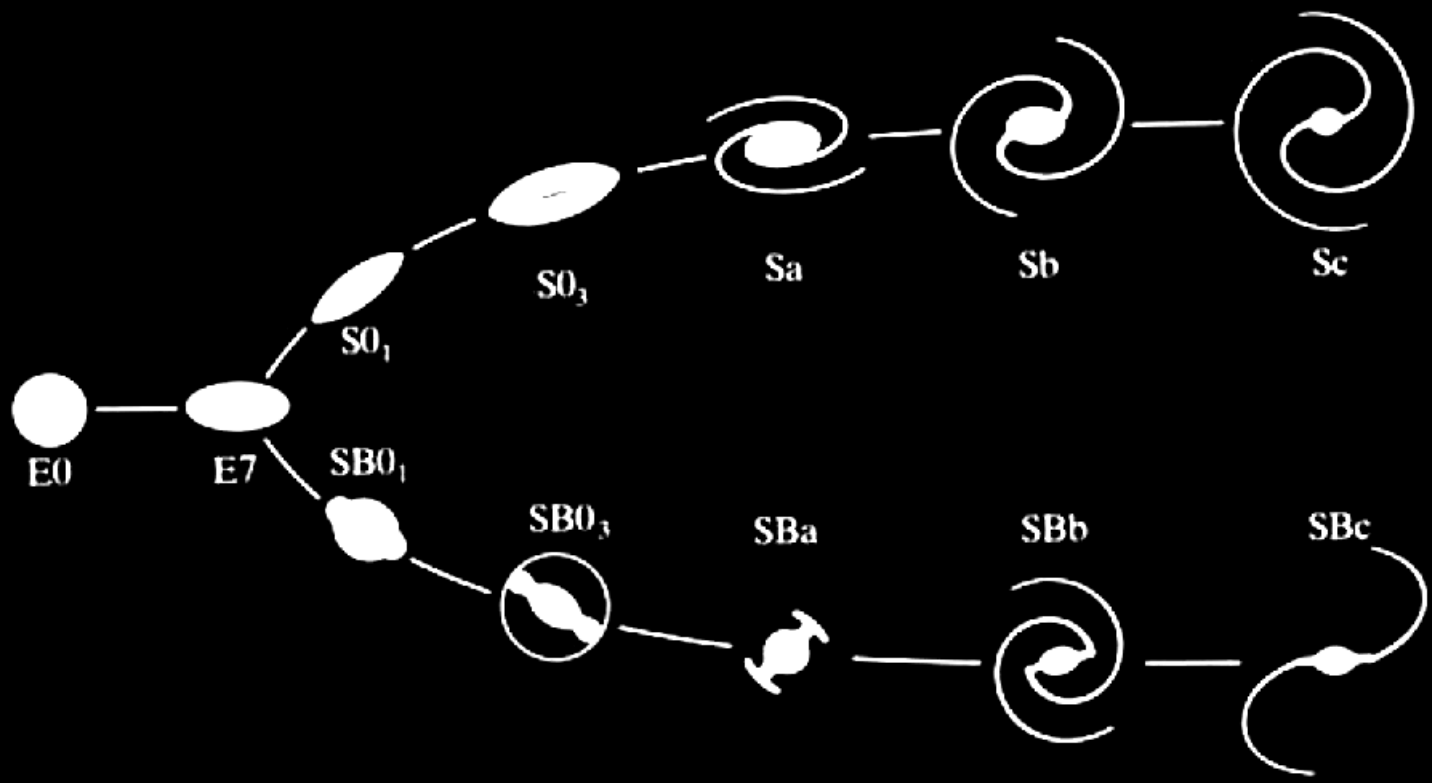
Galaxy Types

Are all galaxies the same?

Galaxy Type	Hubble	de Vaucoulers
Spiral	S, Sa, Sb…	1 through 6
Elliptical	E	-6 through –4
Dwarf	dE, dSph	
Lenticular	S0, SB0	-3, -2, -1
Irregular	Irr	

Hubble's Tuning Fork

Galaxies in The Local Universe



Early Type

Late Type

Spiral Galaxies

- Thin disks
- Most have some form of a bar – arms will emanate from the ends of the bars
- Other classification:
  1. Relative importance of central luminous bulge and disk in overall light from the galaxy
  2. The tightness of the winding of the spiral arms
  3. Degree to which spiral arms are resolved into stars and individual HII regions

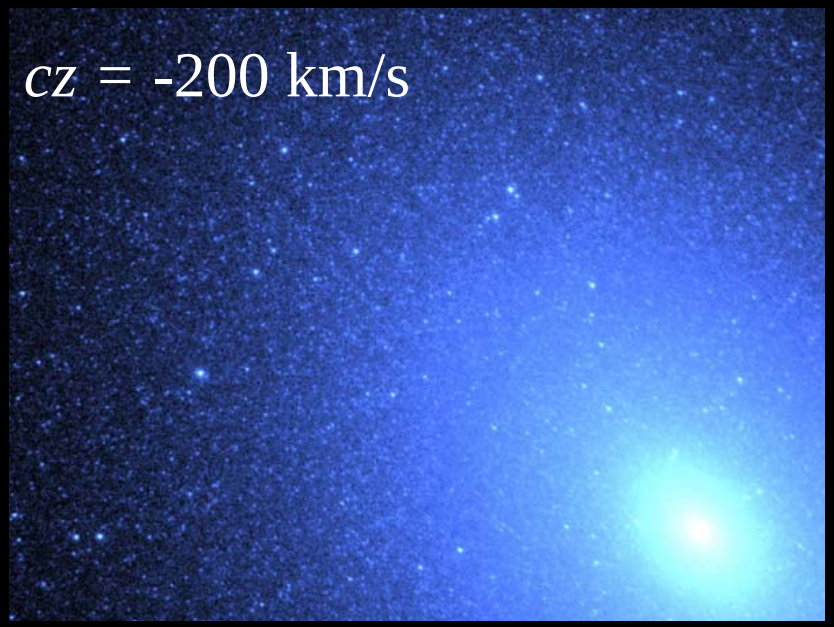
M51
cz = 600 km/s

cz = -179 km/s

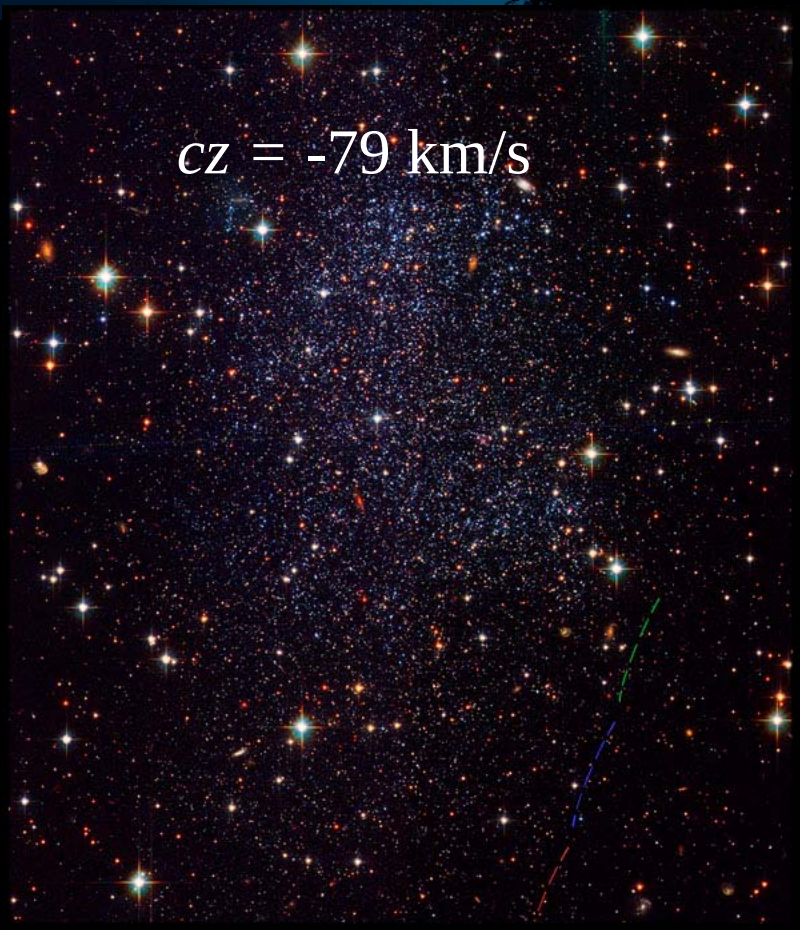
M33

Dwarf Galaxies

- Smaller size than giant elliptical galaxies
- Lower surface brightness
- dE galaxies dominate Virgo



M32



Sagittarius Dwarf

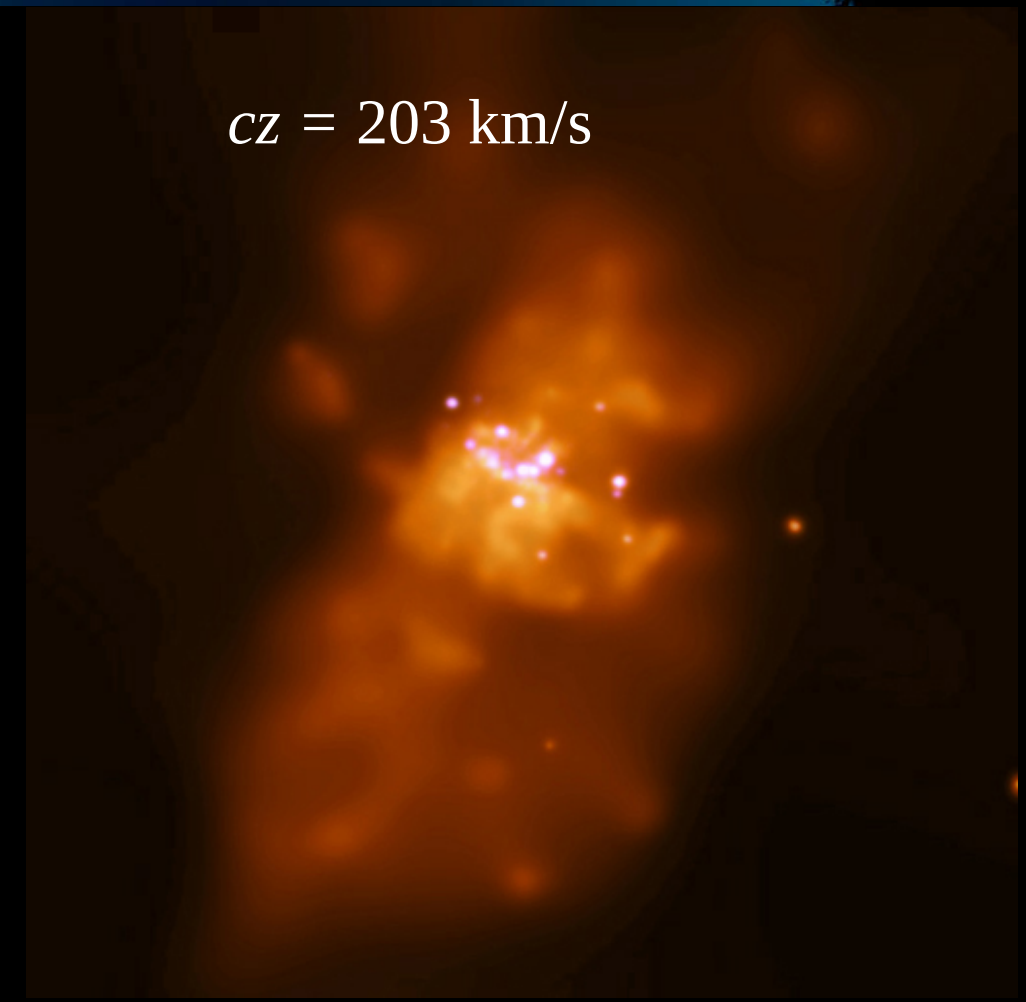
Irregular Galaxies

- LMC and SMC are satellite galaxies of our own – disrupted by gravitational interaction with the Milky Way…

LMC and SMC

Irregular Galaxies

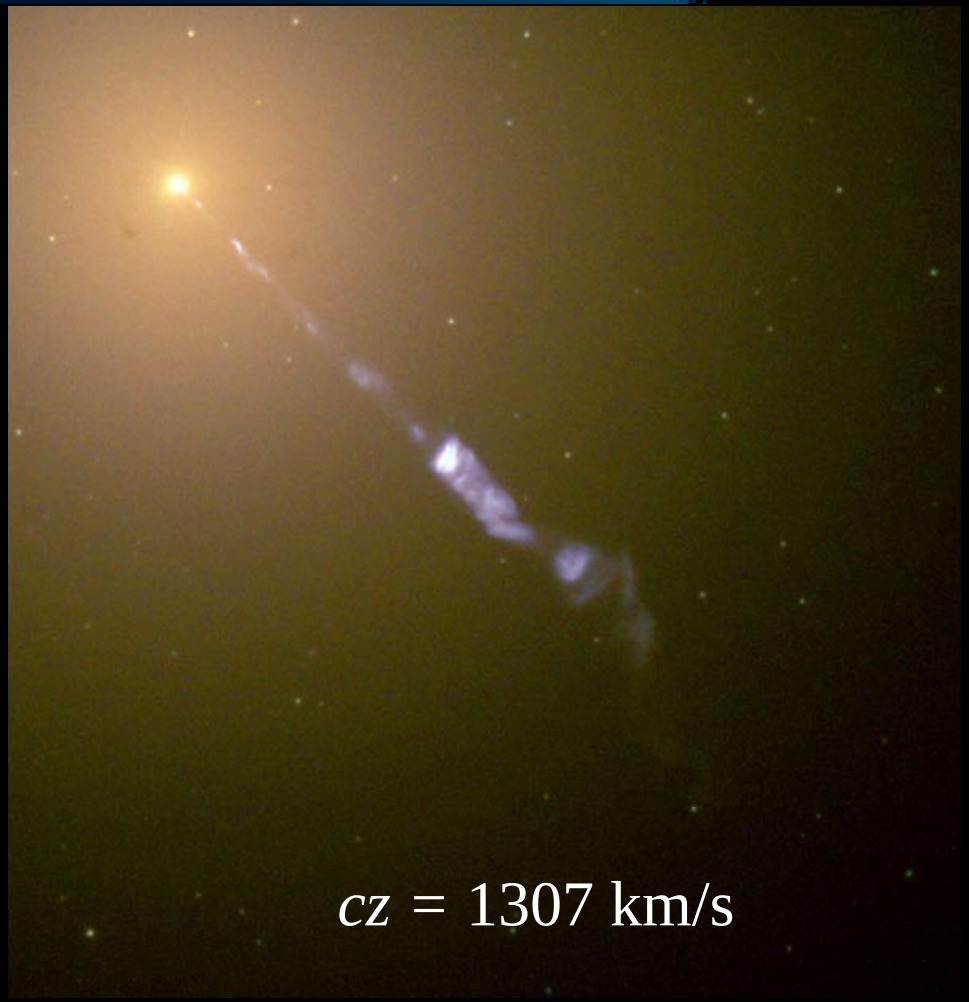
- M82 – irregular starburst galaxy
- Star formation rate at **10 times** the rate of our galaxy
- Chandra X-ray image reveals hot gas flowing out of the galaxy – hot spots indicate x-ray binary stars – some of the brightest known!



Elliptical Galaxies

- Smooth and very little structure; varying in shape
- Classified by E*N* where

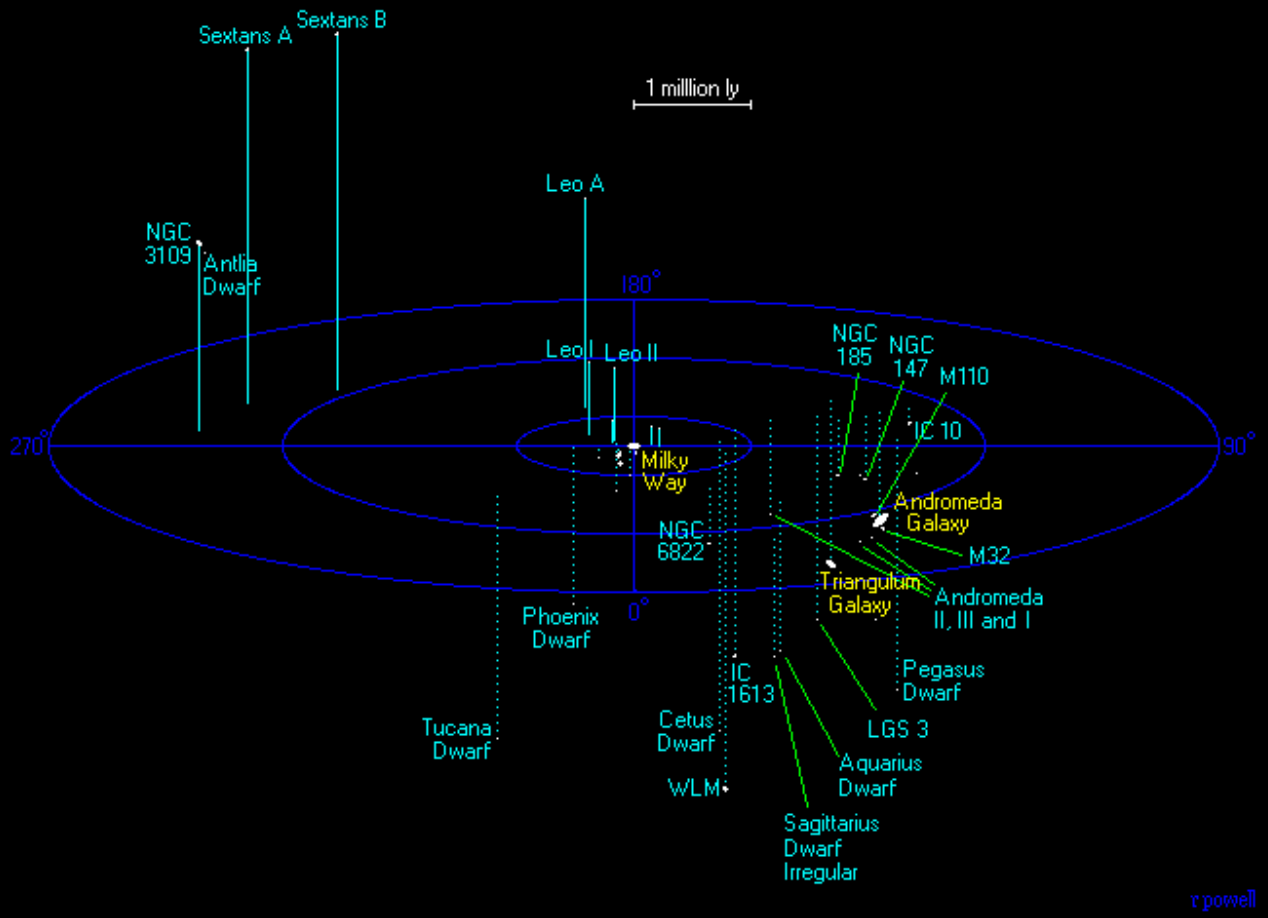
  $N = 10(1 - b/a)$
- Large populations in clusters.
- Little gas – don't see spectral HI lines



M87

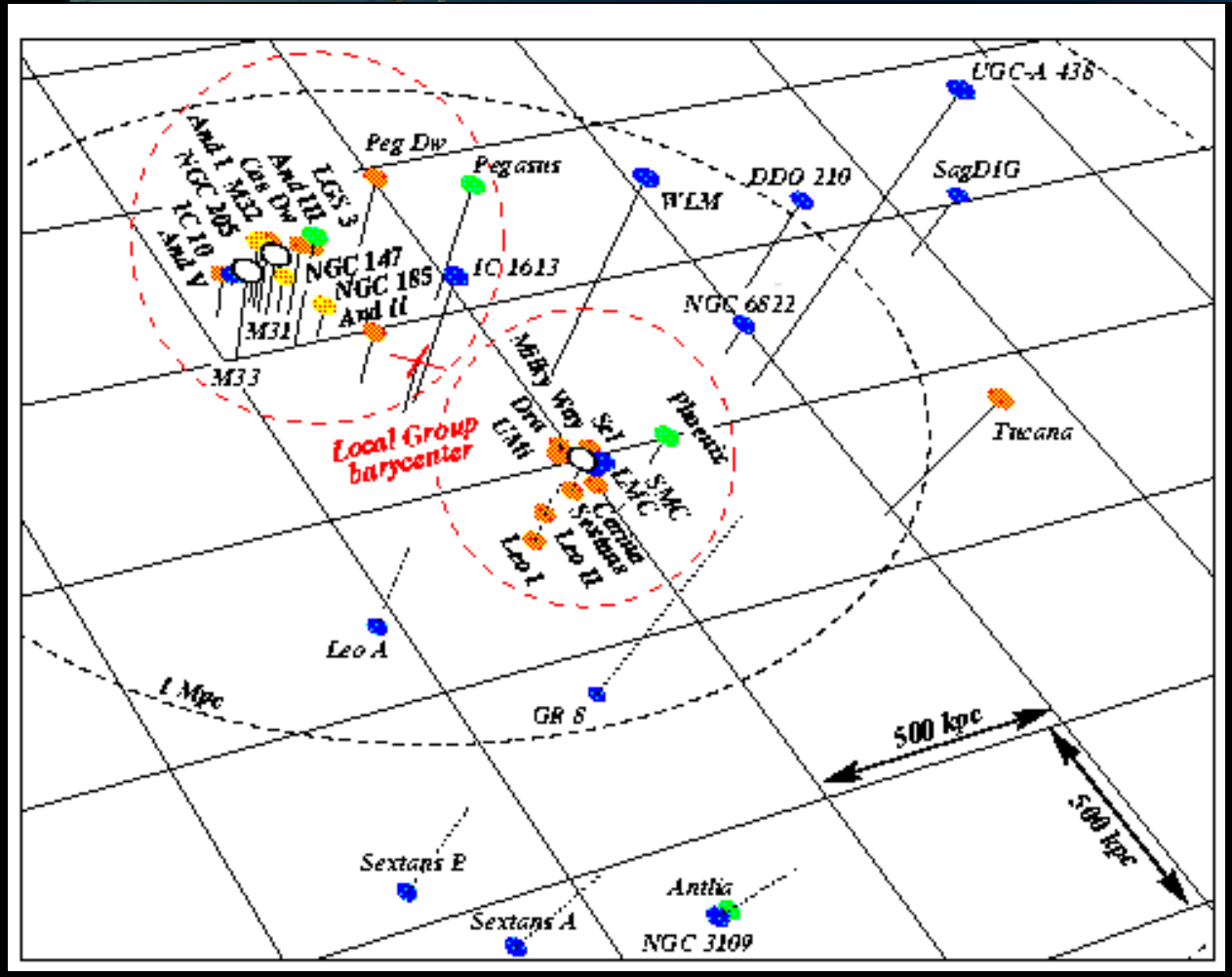
The Local Group

- The Local group has 41 members, ranging from large spiral galaxies to small dwarf irregulars. Most galaxies are dwarf spheriodals…



The Local Group

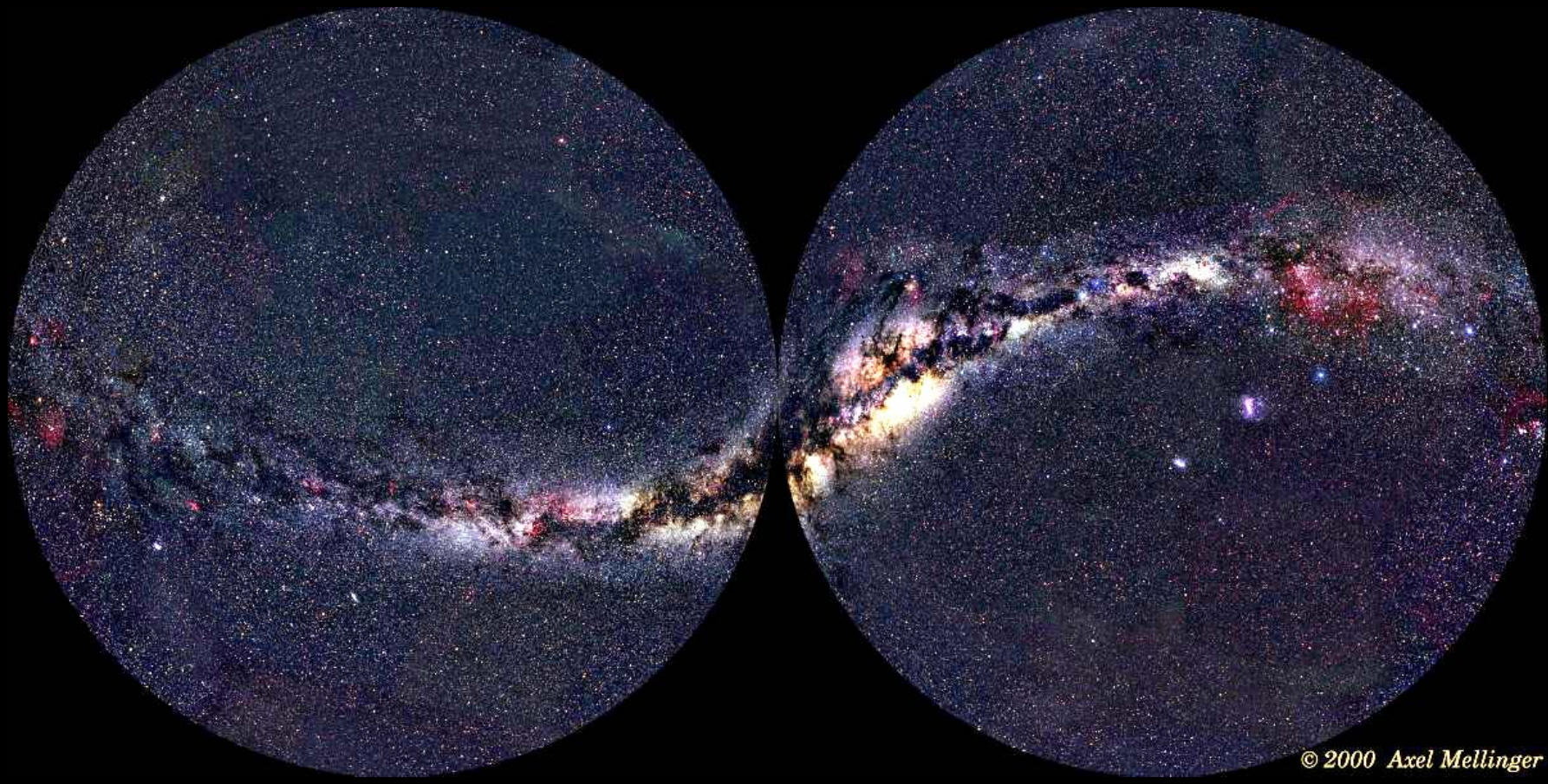
Galaxies in The Local Universe



- Giant spirals
- dSph (+dEll)
- dIrr
- dIrr/dSph

The Milky Way Galaxy

- An Sbc galaxy that is 30 kpc in diameter

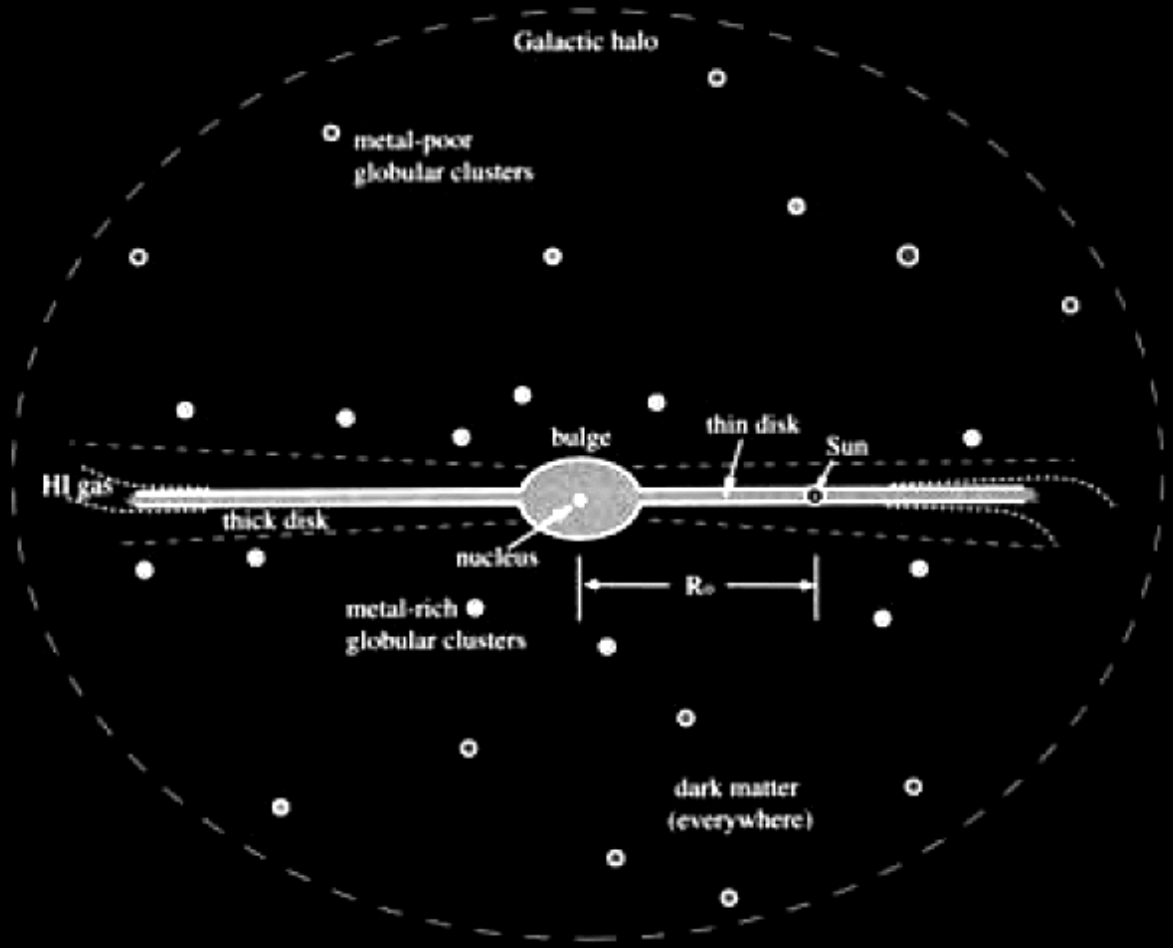


Anatomy of the Milky Way

Galactic component	h_z (pc)	σ_R (km s^{-1})	σ_ϕ (km s^{-1})	σ_z (km s^{-1})	$\langle v_y \rangle$ (km s^{-1})
HI gas near the Sun	130		≈5	≈7	tiny
Local CO, H_2 gas	65		4		tiny
Disk stars: $Z > Z_\odot/4$		(Fig. 2.8)			
$\tau < 3$ Gyr	≈250	30	21	16	−11
$3 < \tau < 6$ Gyr	≈300	36	25	19	−9
$6 < \tau < 10$ Gyr	≈350	38	25	24	−16
$\tau > 10$ Gyr		62	52	37	−21
Thick disk					
[Fe/H] > −0.8	~1500	52	37	40	−35
Halo stars near Sun					
[Fe/H] < −1.6	≳1 kpc	~150	~100	~100	−210
Halo stars at $2.5R_0$	few kpc	80–100	130–150	130–150	−220

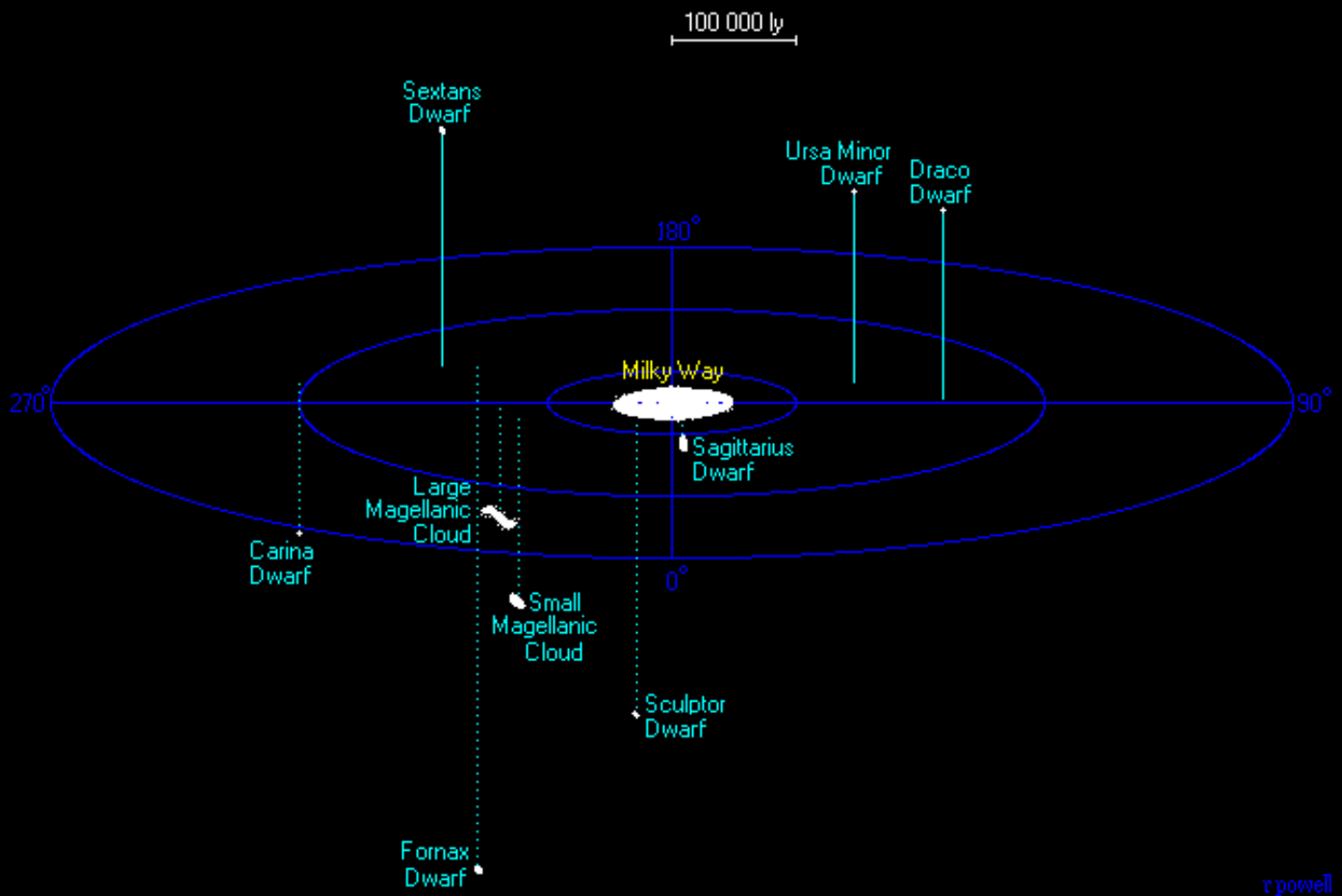
Anatomy of the Milky Way

- R_0 ~ 8 kpc
- 200 billion stars
- M_{tot} 5 x 10^{11} $M_\odot$
- SFR ~ 3 $M_\odot$/yr
- Bulge ~ 3 kpc in diameter

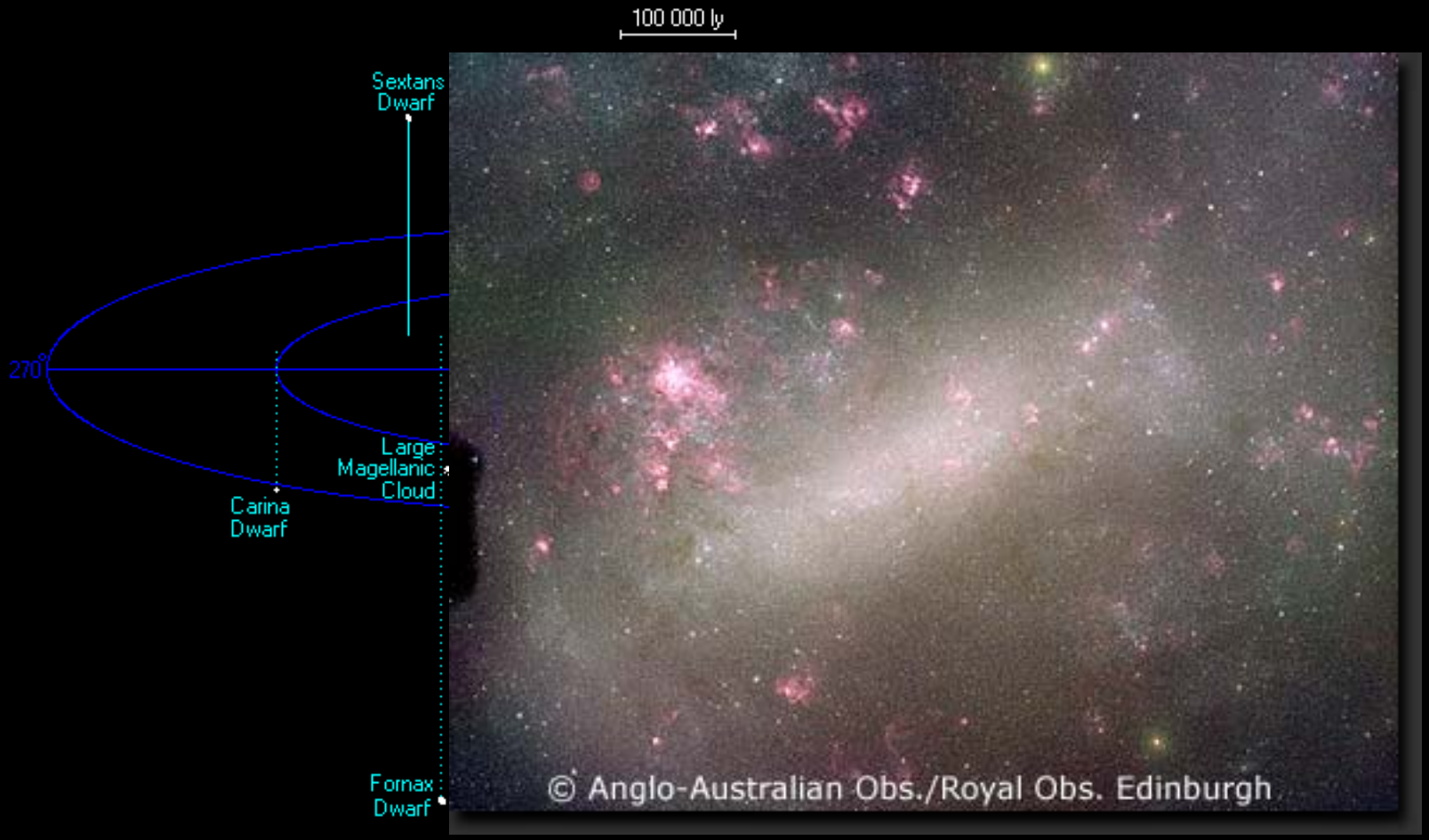


Around the Milky Way…

Galaxies in The Local Universe



Around the Milky Way…



Around the Milky Way…

The Andromeda Galaxy

- Sb galaxy 770 kpc from the Milky Way.
- Larger, more luminous, with a larger disk scale length than the Milky Way – it even rotates faster at 260 km/s!
- At least 9 known satellite galaxies – dwarf elliptical and spheroidals!

cz = -300 km/s

The Andromeda Galaxy

- GALEX mission mosaic in the ultraviolet

cz = -300 km/s

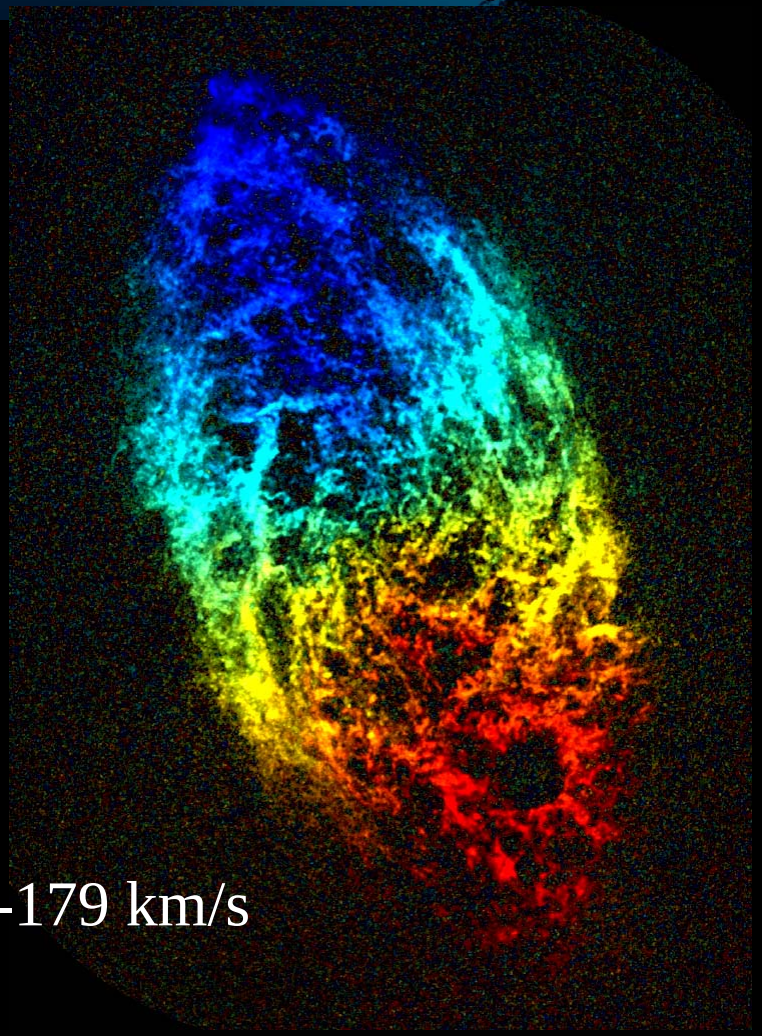
M33

- Late-type spiral galaxy ~850 kpc from the Milky Way and ~200 kpc from Andromeda
- Disk scale length is around 1.7 kpc, rotating around 120 km/s.

cz = -179 km/s

M33

- Richer in HI gas than M31 or the Milky Way – VLA doppler image show movement of the HI gas towards and away.
- The HI disk extends out to 30 kpc, enough for M31 to cause tidal effects and warp the outer disk!



cz = -179 km/s


Galaxy Groups

Groups of galaxies

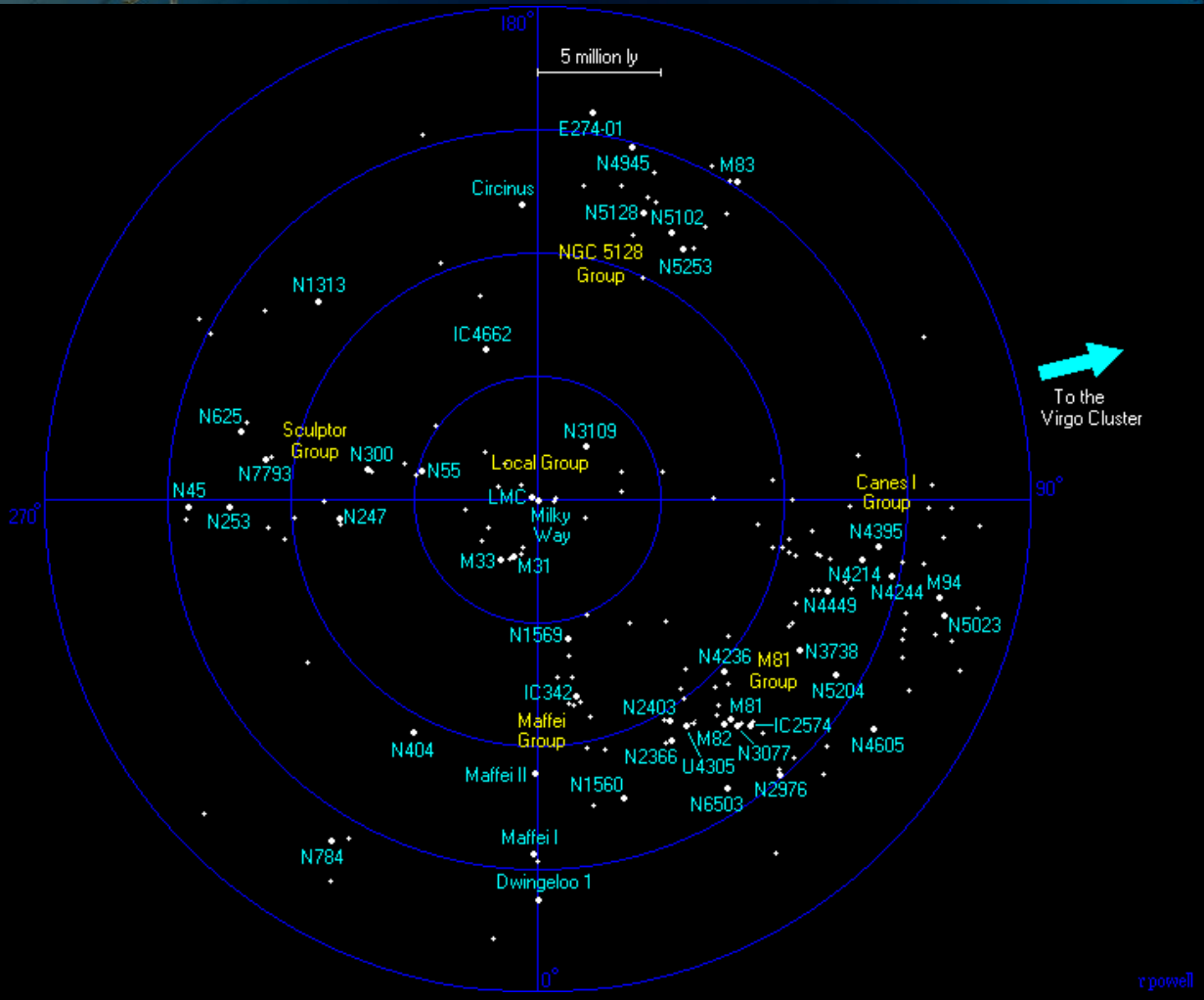
- Galaxies can be gravitationally bound to each other, and undergo interactions and collisions.
- Separations across intergalactic distances range from 50 kpc up to 1 Mpc.
- Groups are important because one can determine a dynamical mass for the system.
- ALFALFA science goals include studying the effects within the group environment –
  - What is HI mass function?
  - How do unseen HI clouds/starless galaxies effect dynamics?
  - Are their unseen tidal remnants or debris?
  - What are sizes of HI disks?

Neighboring Galaxy Groups

Parameter	Milky Way	M31	M81	Cen A	M83	IC 342	Maffei	Sculptor[a]	CVn I[a]
D_{MW} (Mpc)	0.01	0.77	3.63	3.66	4.56	3.28	3.01:	3.94	4.09
D_{LG} (Mpc)	0.43	0.34	3.47	4.10	4.98	2.94	2.67:	3.79	4.17
SGZ[b] (Mpc)	0.00	0.07	0.04	−0.33	0.08	0.02	0.08	−0.34	0.77
N_{tot}	15	19	29	28	14	8	8:	6	9
N_{E+dSph}	10	13	11	18	4	0	1:	3	1
Type(1)	4	3	3	−2	5	5	4	5	2
M_B(1) (mag)	−20.80	−21.58	−21.06	−20.77	−20.43	−20.69	−20.15	−21.37	−19.83
V_m(1) (km s^{-1})	220	255	232	398	211	162	163	199	164
V_{LG}(1) (km s^{-1})	−88	−35	107	301	304	245	212	274	353
$\langle V_{LG} \rangle$ (km s^{-1})	−79	−16	193	312	308	229	302	279	306
σ_v (km s^{-1})	76	77	91	105	71	54	59	54	56
$\langle R_p \rangle$ (kpc)	155	254	211	290	164	322	104	359	385
L_B (10^{10} $L_\odot$)	3.28	6.83	6.11	5.55	2.31	3.21	2.69	5.58	2.00
M_{vir} (10^{10} $M_\odot$)	93	57	117	489	109	57	65	332	267
M_{orb} (10^{10} $M_\odot$)	96	111	197	288	100	95	135	153	322
M_{vir}/L ($M_\odot$ $L_\odot^{-1}$)	28	8	19	88	47	18	24	60	133
M_{orb}/L ($M_\odot$ $L_\odot^{-1}$)	29	16	32	52	43	30	50	28	161
T_{cross} (Gyr)	2.1	3.3	2.3	2.8	2.3	5.9	1.8	6.6	6.9

Neighboring Galaxy Groups

Galaxies in The Local Universe

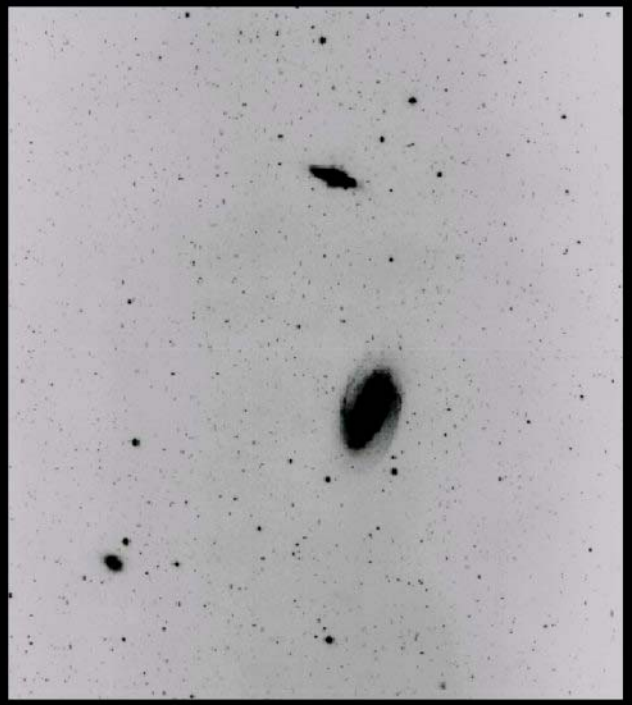


Groups of galaxies

TIDAL INTERACTIONS IN M81 GROUP

Stellar Light Distribution

21 cm HI Distribution

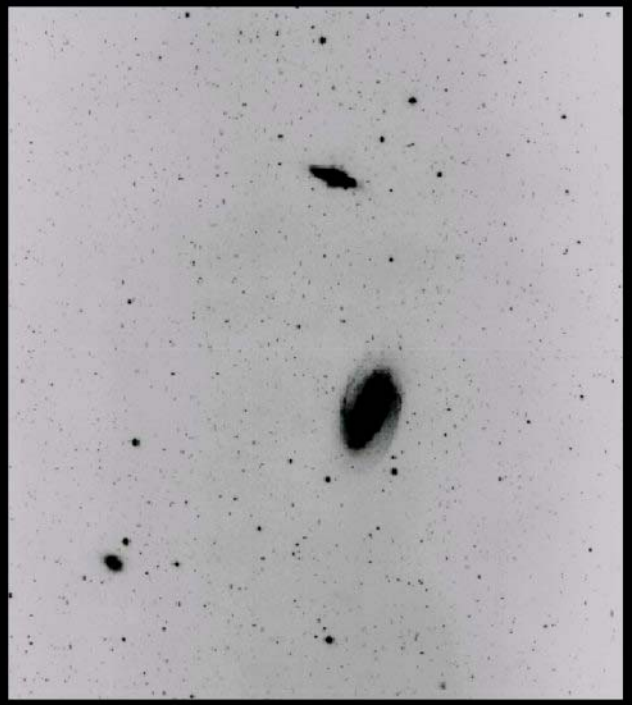


cz = -34 km/s

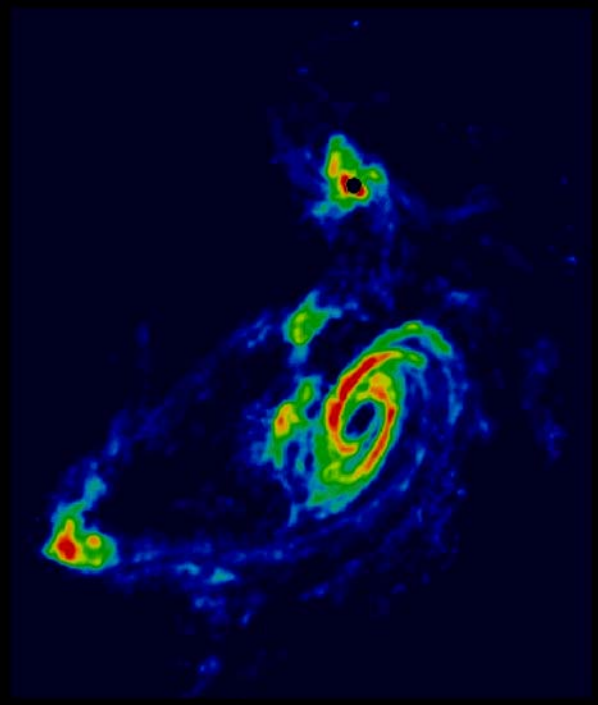
Groups of galaxies

TIDAL INTERACTIONS IN M81 GROUP

Stellar Light Distribution



21 cm HI Distribution

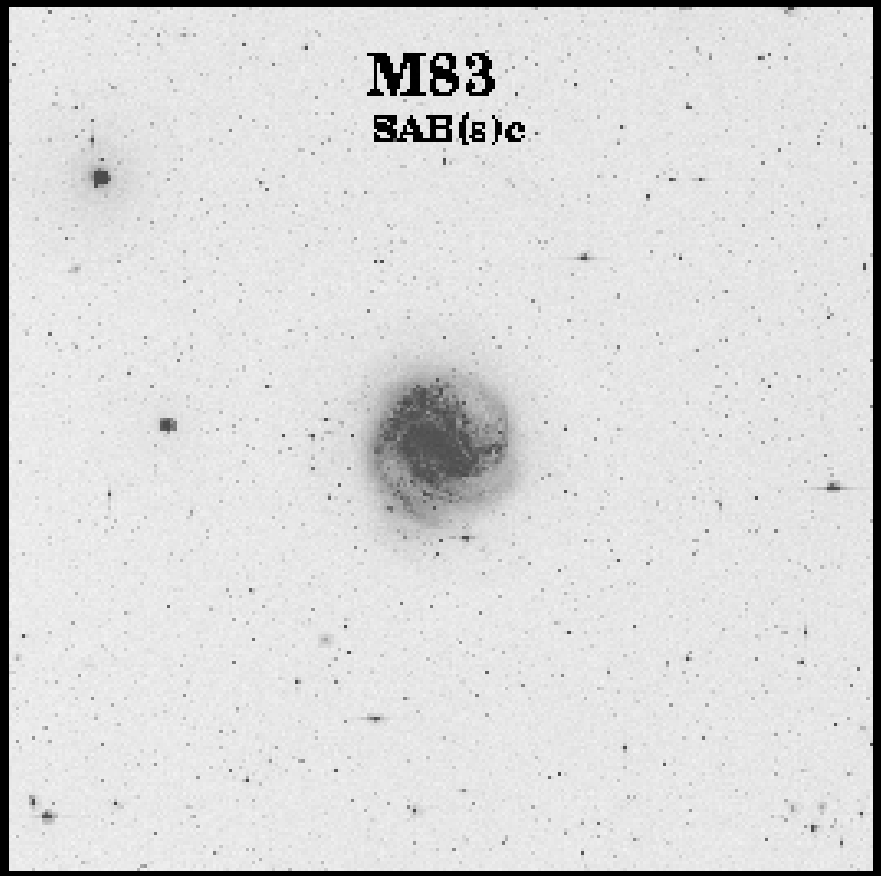


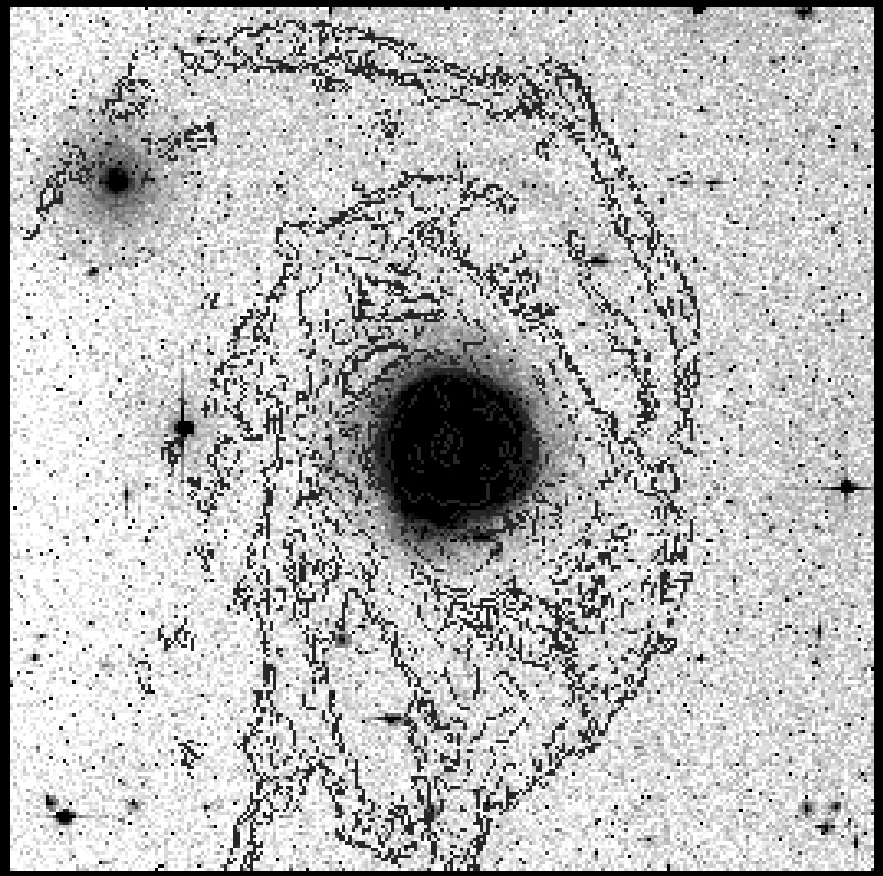
cz = -34 km/s

Cen A Group

cz = 547 km/s

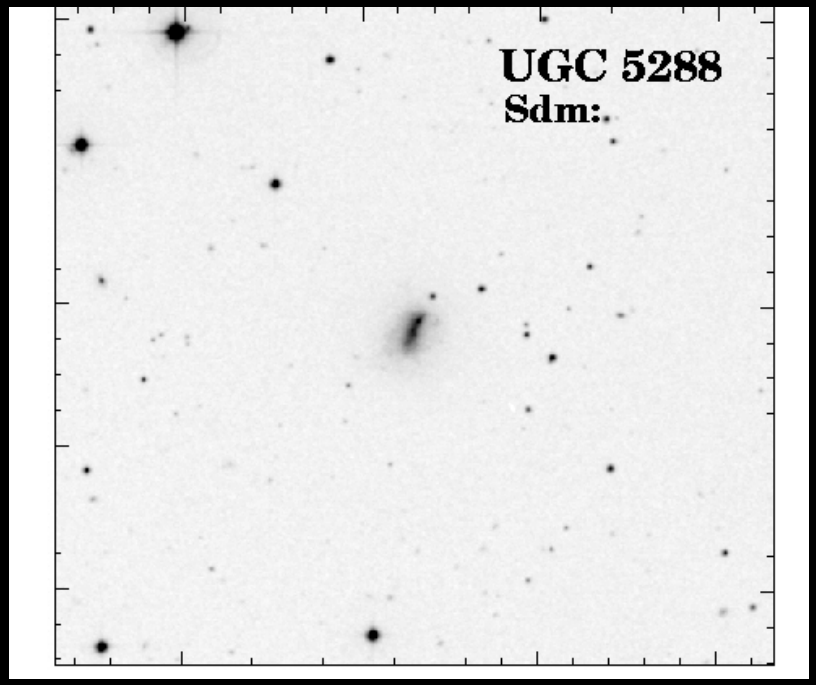
M83 Group

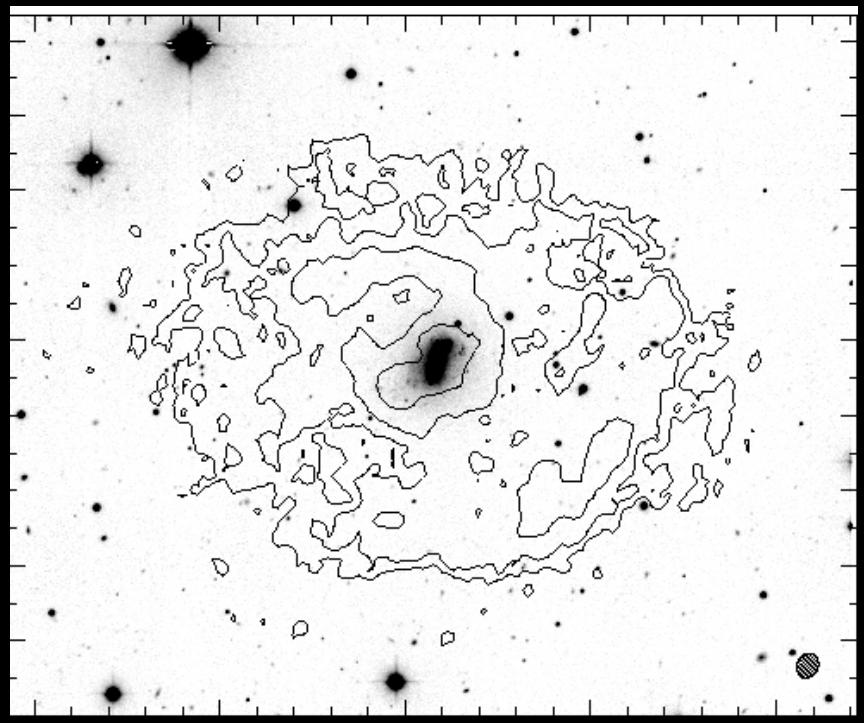




cz = +513 km/s

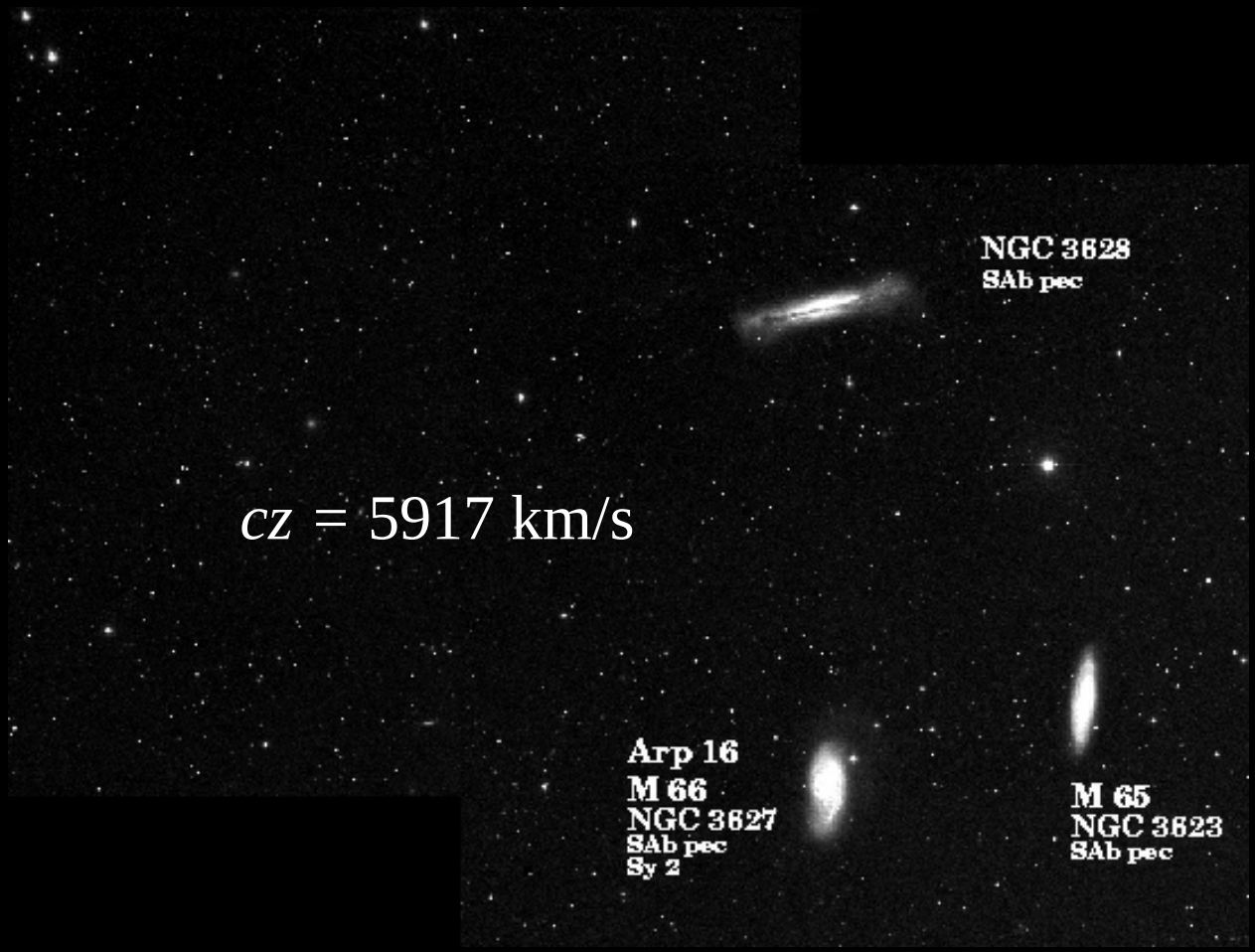
Extended HI Disk of a BCD



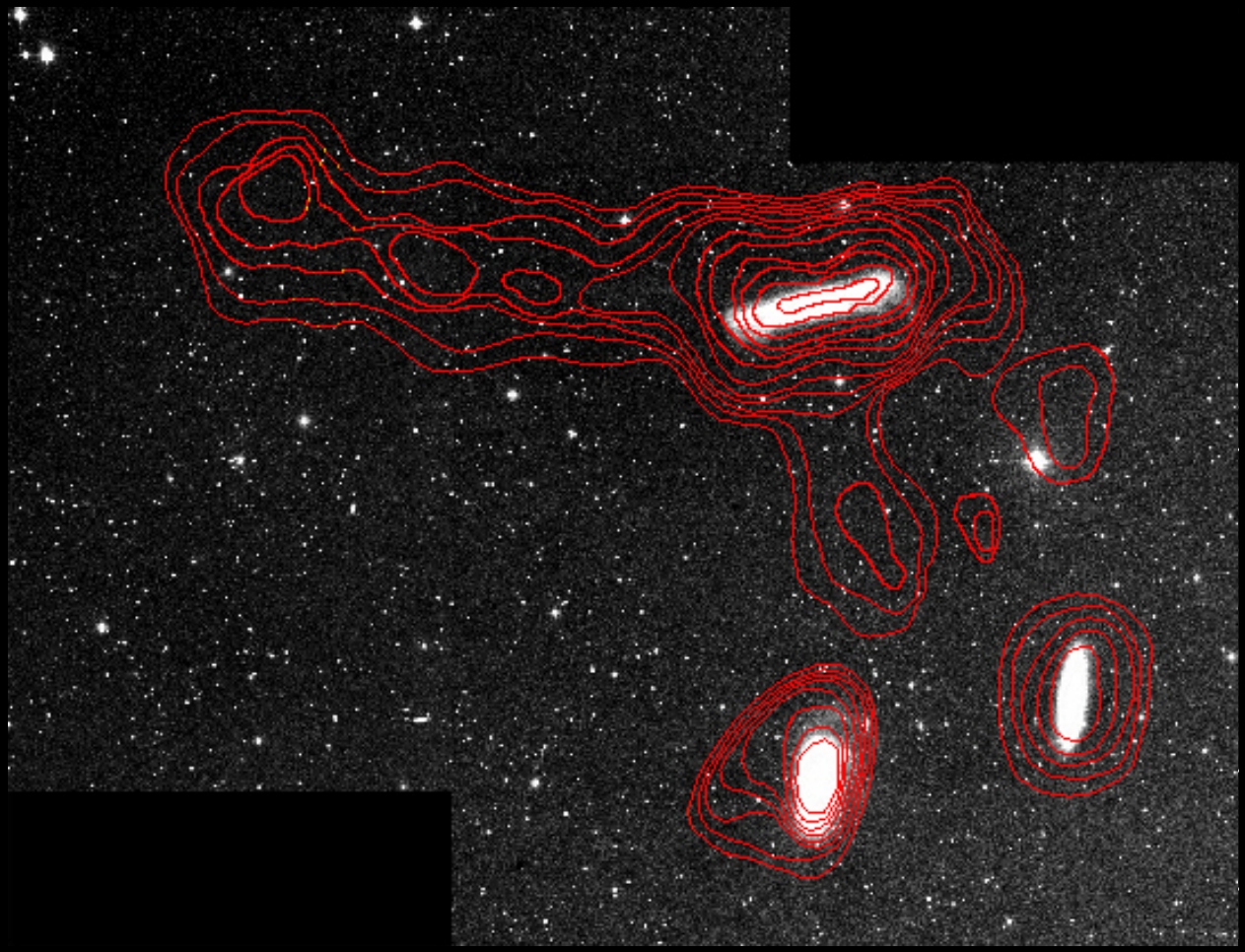


cz = 557 km/s

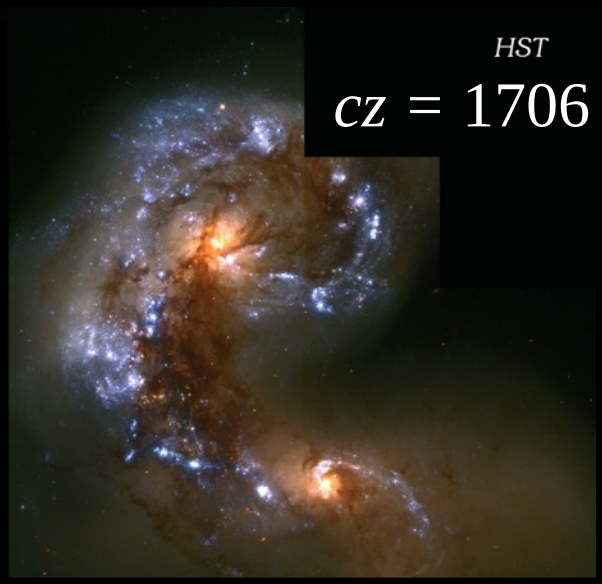
M66 Group


NGC 3628
SAb pec
cz = 5917 km/s
Arp 16
M 66
NGC 3627
SAb pec
Sy 2
M 65
NGC 3623
SAb pec

M66 Group



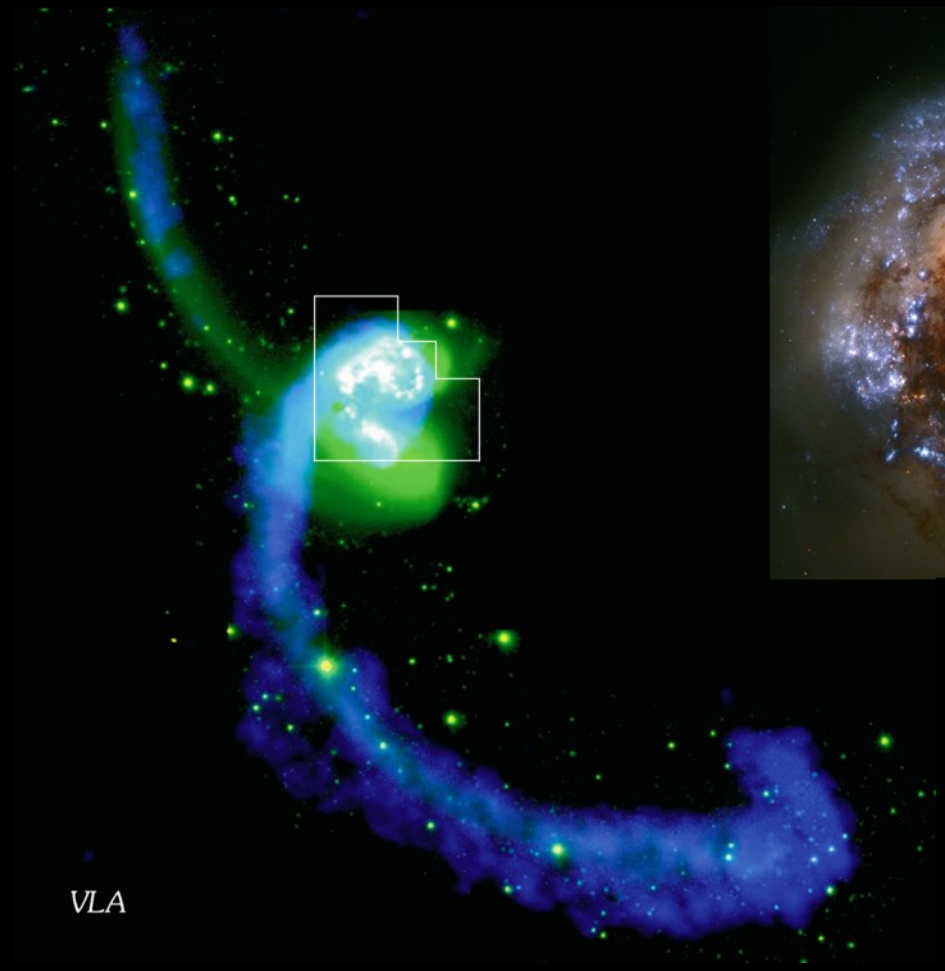
Colliding Galaxies

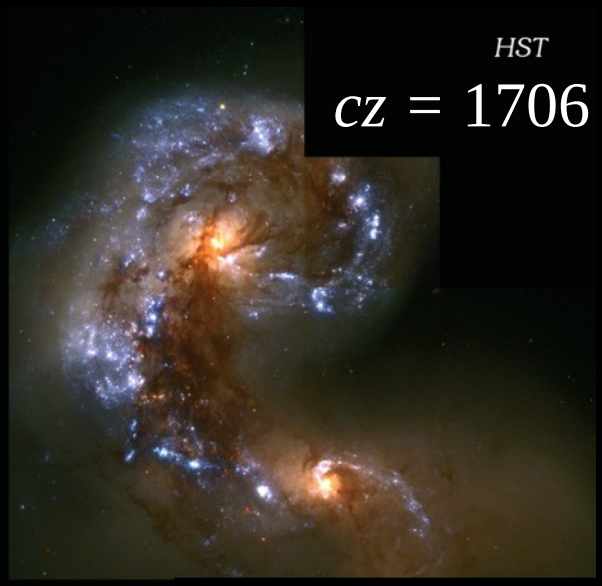


HST

cz = 1706 km/s

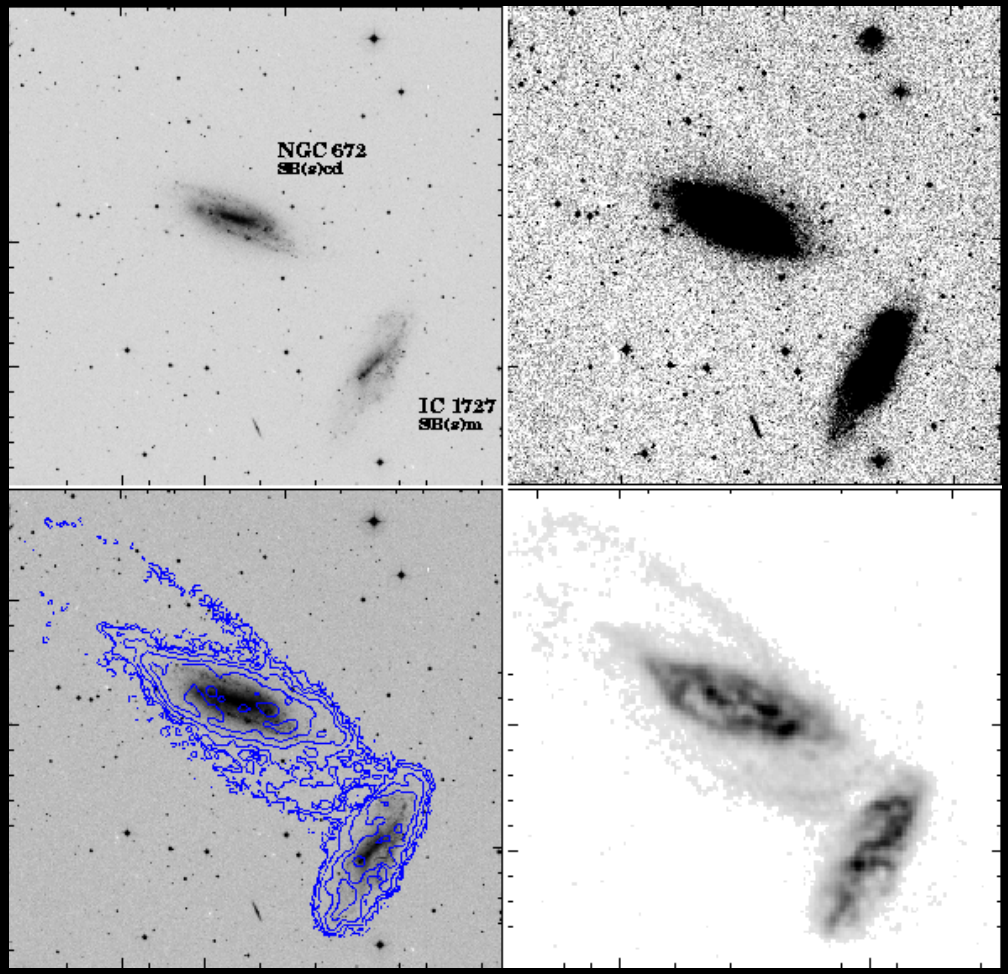
Colliding Galaxies





cz = 1706 km/s

Colliding Galaxies



cz = 422 km/s


Galaxy Clusters

Clusters of Galaxies

- Around half the galaxies in the Universe are found in clusters or groups.
- Cluster have a higher density than “loose” groups – brightest galaxies are S0s and ellipticals instead of spirals
- Abell Catalog contains 4073 rich clusters
- Gravity binds the members, as well as hot intracluster gas (seen in the X-ray)

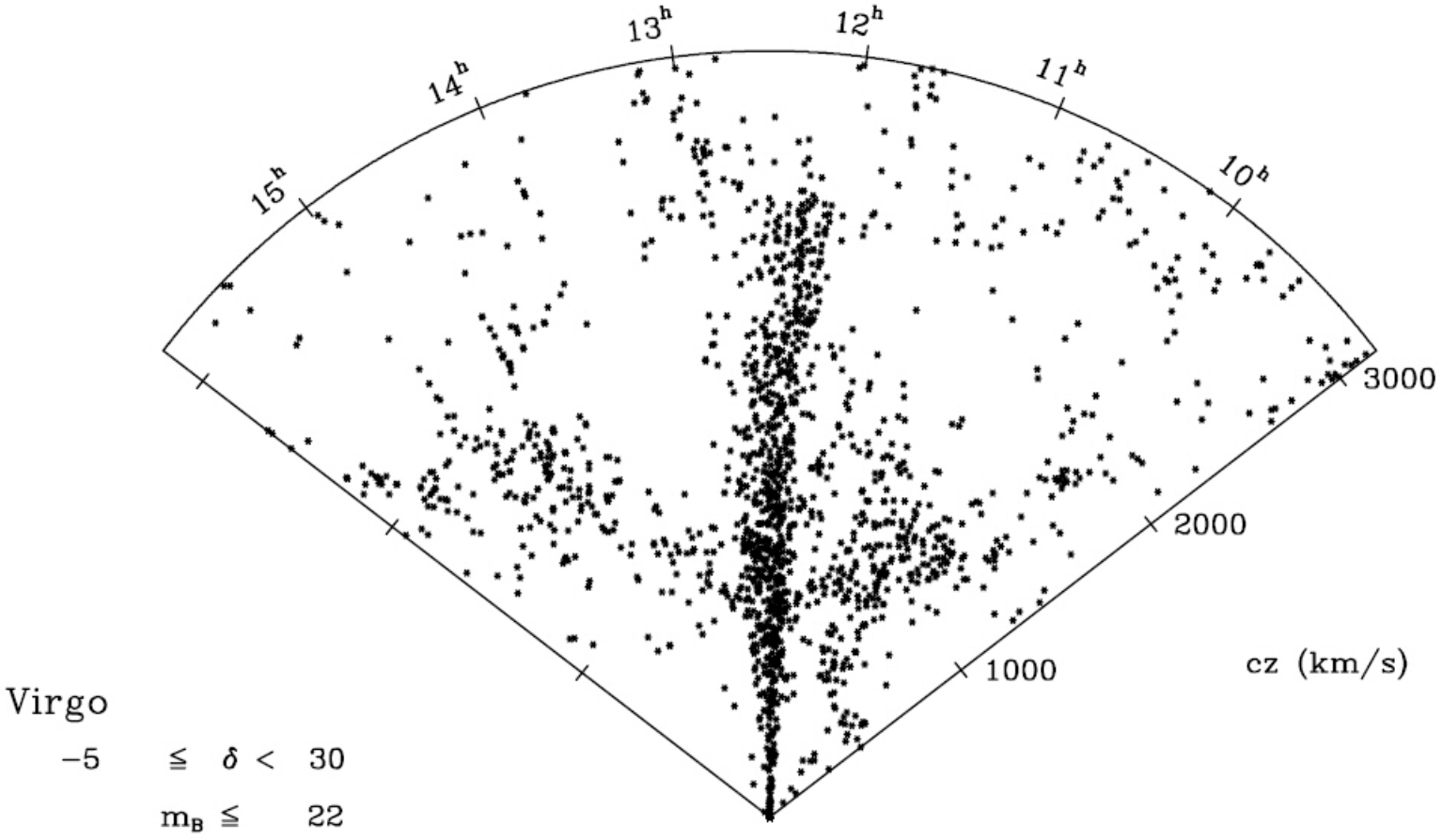
Clusters of Galaxies

- Some well known clusters:
  - Virgo
  - Fornax
  - Eridanus
  - Coma
  - Perseus
  - Hercules
  - Leo
  - Centaurus

Virgo Cluster

- cz ~ 1035 km/s
- Δv ~ 1000 km/s !!
- 1300 catalogued members!!
- Most galaxies are dwarf elliptical type
- Core radius ~ 500 kpc

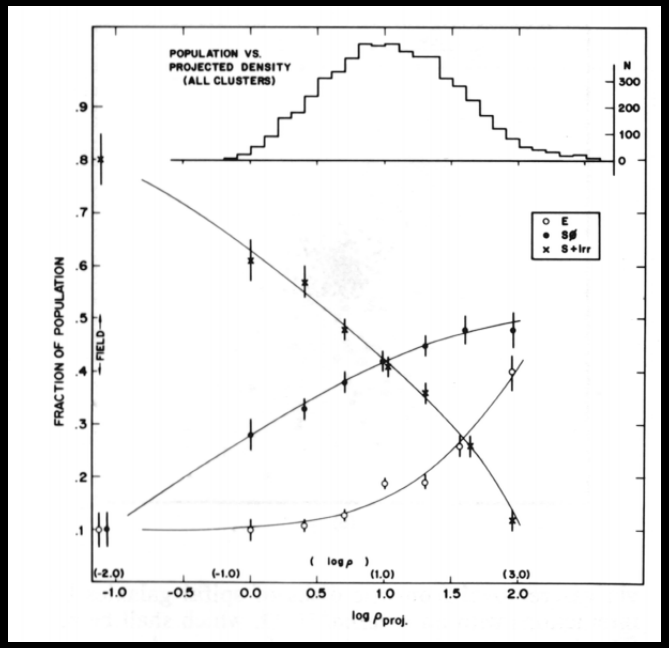
Virgo Substructure



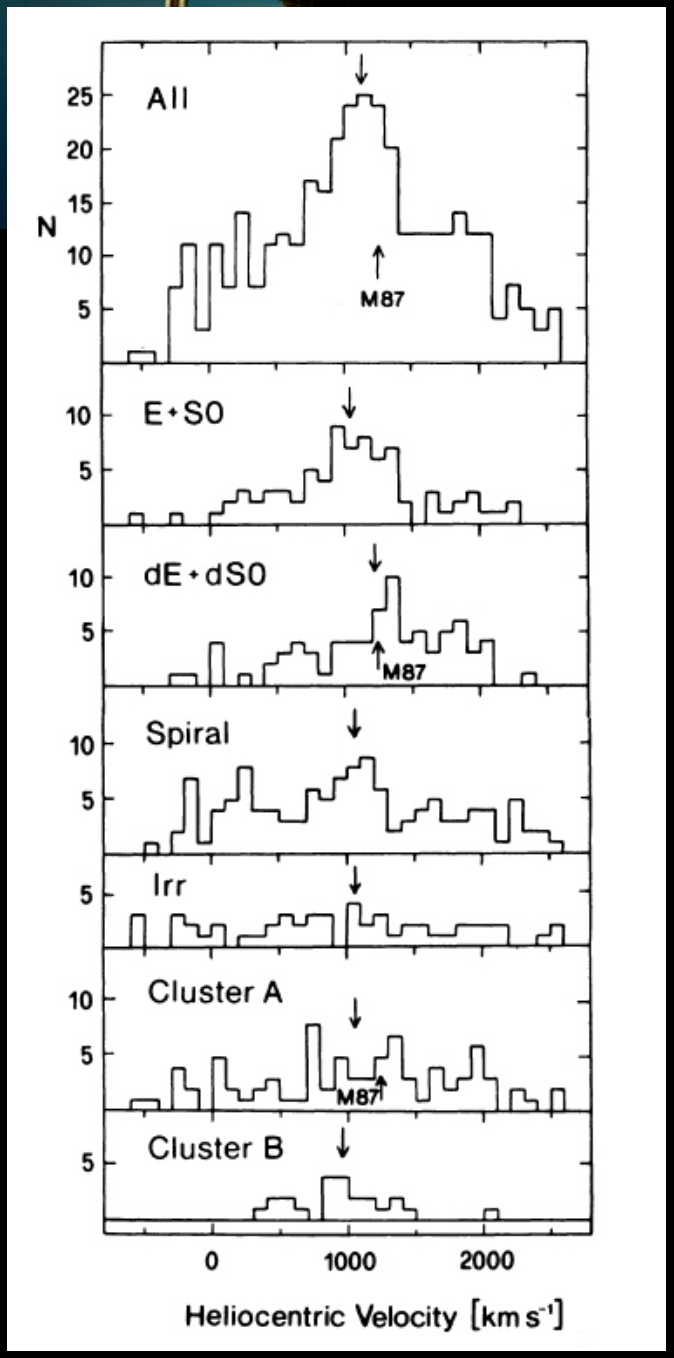
2001

Galaxy distributions

- The E+S0 galaxy distribution are more gaussian in nature.
- The wider dispersion in the late-type galaxies is indicative of the infall/expansion regions
- Morphology – density relation details that the spiral fraction decreases in higher density regions, such as that of a cluster.



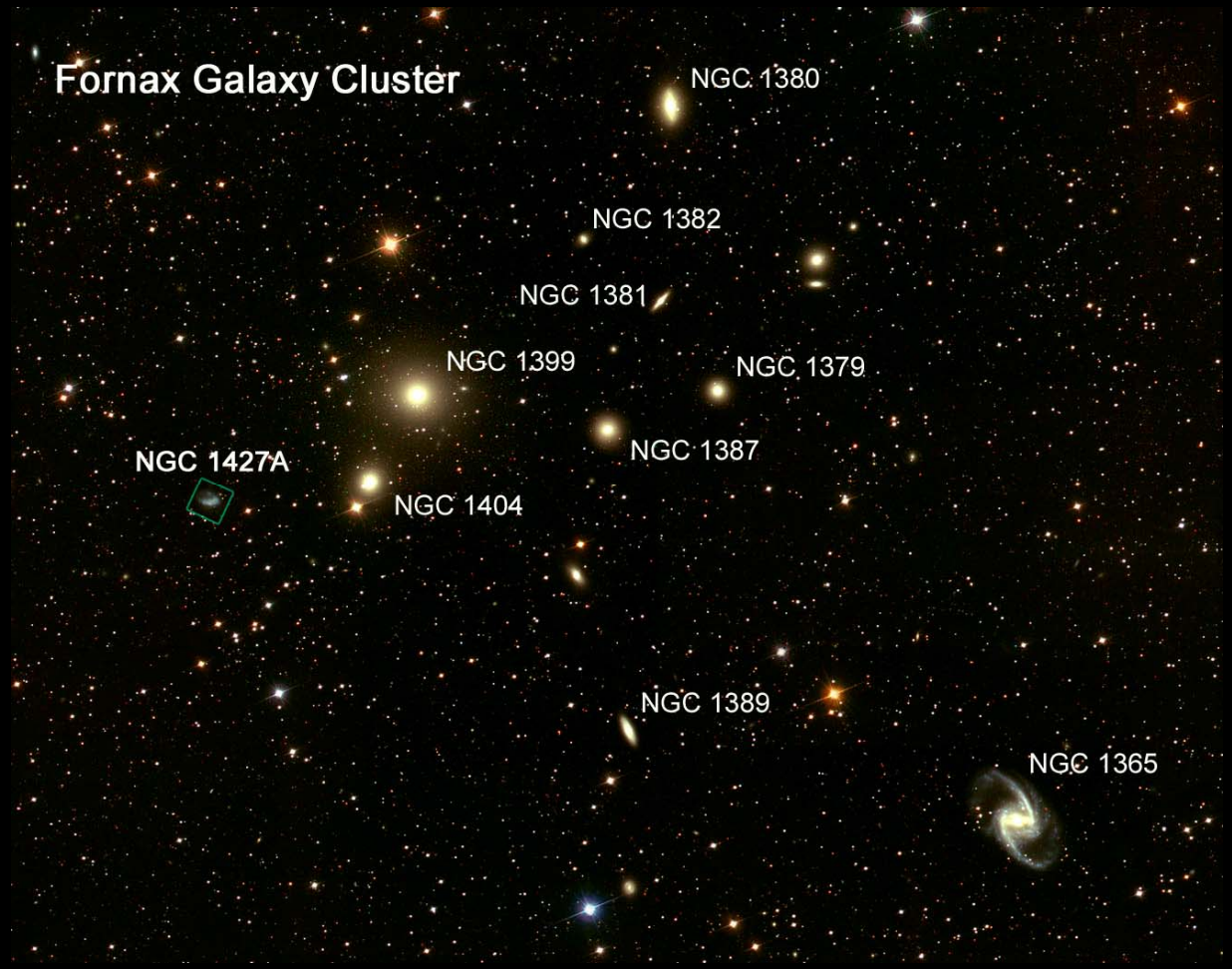
Dressler 1984



Binggeli et al. 1987

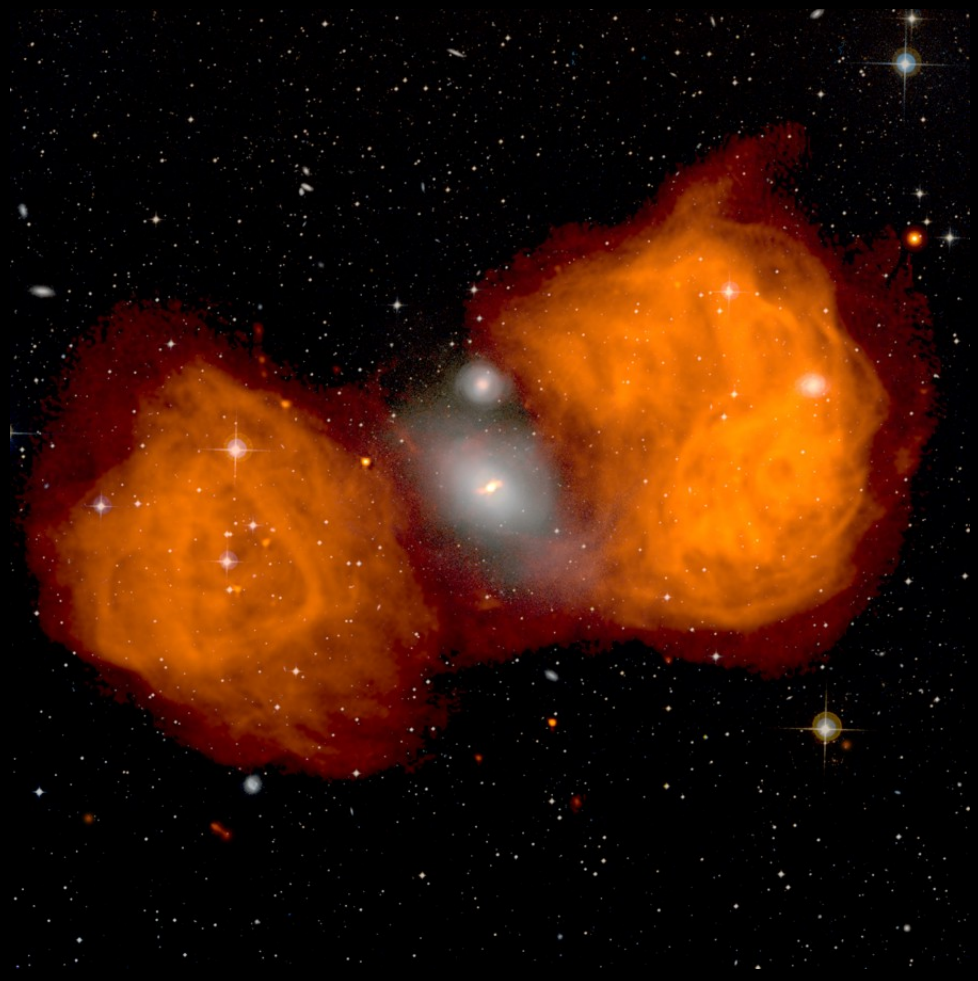
Fornax cluster

- $cz \sim 1400$ km/s



Fornax Cluster – Xray view

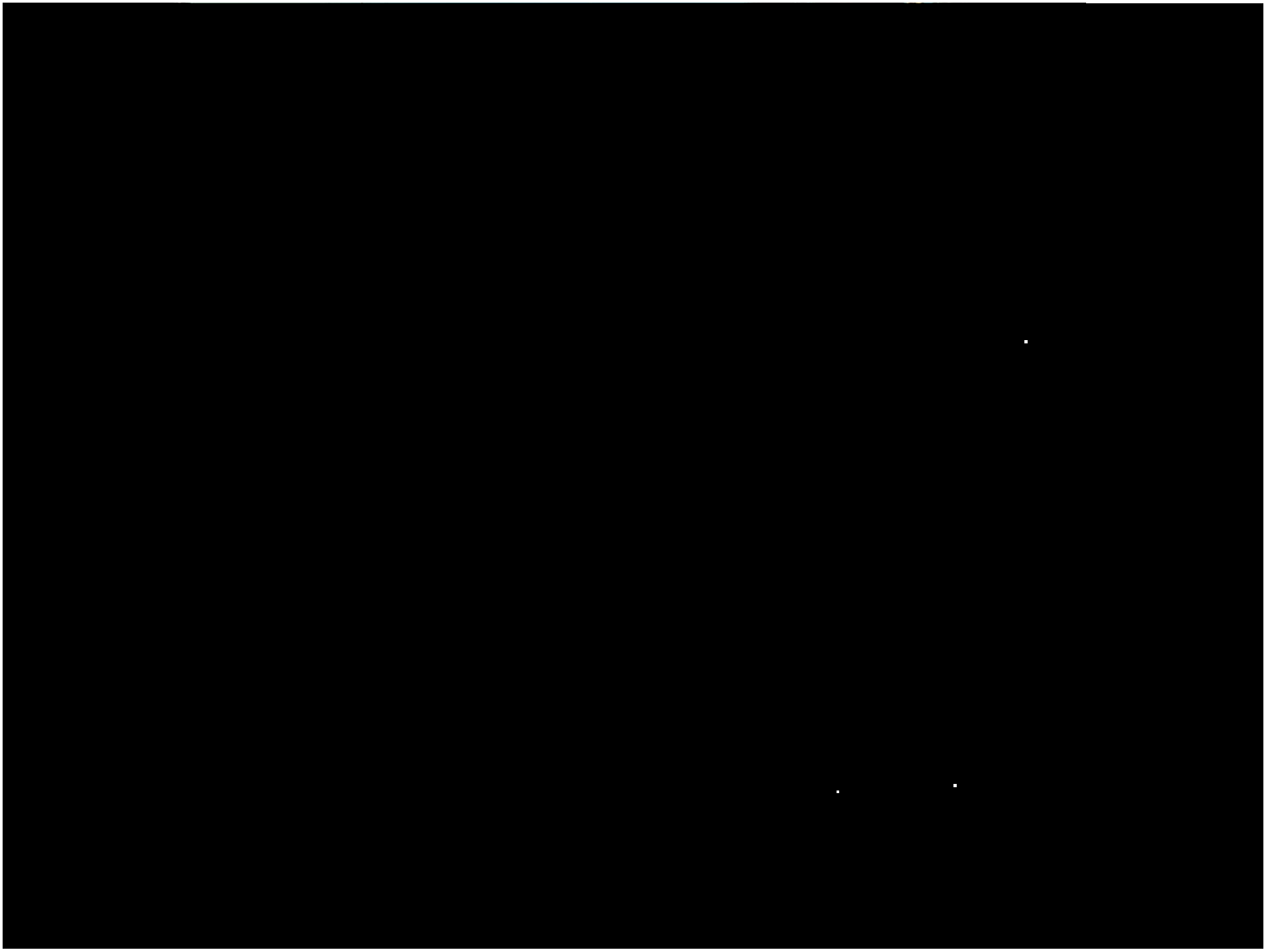
Fornax – optical + radio



View of the Local Universe

Local Universe Overview

- Springob, C.M. et al. 2005, ApJS. (in press)
- ~9000 redshifts based on HI detections
- Data taken from a variety of radio telescopes…



References and Acknowledgements

- Hibbard's Rogue's Gallery
- Herter, T. Astro 530, Cornell University
- Malin, D. Anglo-Australiian
- Alex Mellinger
- Mateo, et. al.
- APOD
- Haynes, M. P.
- Giovanelli, R.
- J. Sanders, A. Fabian
- KPNO REU Program
- Nigel Sharp, KPNO, WIYN
- Jason Ware
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